

*Letter to the Editor***Evidences for two Gamma-Ray Burst afterglow emission regimes****M. Boër\* and B. Gendre**

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**Abstract.** We applied cosmological and absorption corrections to the X-ray and optical afterglow fluxes of a sample of Gamma-Ray Burst sources of known distance. We find a good correlation in X-rays and that the GRBs in our sample form two well defined classes. We tentatively interpret them as radiative and adiabatic afterglow behaviours in the framework of the fireball model for GRBs. We do not observe this correlation at optical wavelengths. This discrepancy with the model may be due to the absorption in the source vicinity.

**Key words:** gamma rays: bursts**1. Introduction**

The detection of X-ray and optical afterglows of cosmic gamma-ray bursts (hereafter GRB) has firmly established the fireball model (Rees and Mészáros 1992, Mészáros and Rees 1997, Panaitescu et al. 1998) as a standard tool to interpret GRB afterglows. A vast majority of sources have X-ray afterglows, while about half of them have been observed at optical wavelengths. The fireball model provides firm predictions on the temporal behaviour of the afterglow emission at all wavelengths, allowing an inter-comparison of different events (Piran 1999, Sari et al. 1998). In this framework the afterglow emission is described as synchrotron emission of accelerated electrons during the shock of an ultra-relativistic shell with the external medium.

In this letter, we use a set of GRB sources detected during their afterglow both at X-ray and optical wavelengths, and for which a firm measure of the source distance has been established. We apply several distance corrections and we take into account the galactic absorption for the optical data. The light-curve of each source of our sample is computed for a standard distance corresponding to a redshift of 1.

**2. The burst sample**

We used a sample of eight GRBs detected both at X-ray, by the NFI instrument on board the BeppoSAX satellite, and optical wavelengths. A redshift measure is available for each of the

**Table 1.** The GRB source sample

Source	Redshift	References
GRB 970228	0.695	1, 2
GRB 970508	0.835	1, 3
GRB 971214	3.42	1, 4
GRB 980613	1.096	1, 5, 6, 7
GRB 980703	0.966	8
GRB 990123	0.61	9
GRB 990510	1.62	10, 11
GRB 991216	1.02	11, 12

- 1 Costa 1999
- 2 Galama et al. 1997
- 3 Pedersen et al. 1998
- 4 Diercks et al. 1998
- 5 Djogovsky et al. 1998
- 6 Hjort et al. 1998
- 7 Halpern et al. 1998
- 8 Vreeswijk et al. 1999
- 9 Galama et al. 1999
- 10 Staneck et al. 1999
- 11 Piro L. 2000
- 12 Halpern et al. 2000

sources of our sample. We did not use the data from GRB 980425 since its association with SN 1998bw remains questionable. The redshifts and the main references on the sources of our sample have been summarized on Table 1.

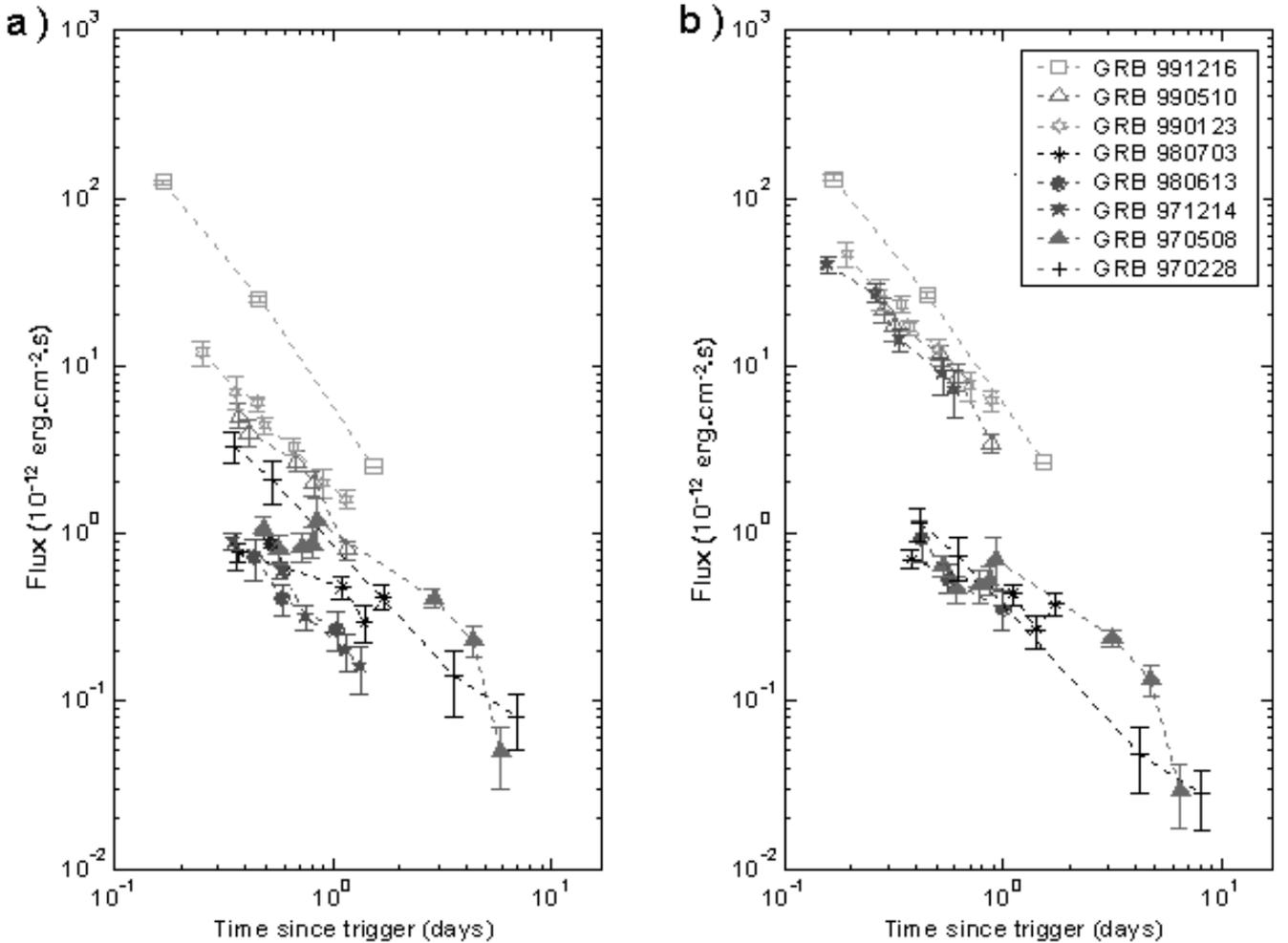
We insist on the fact that we selected GRBs on the basis of the consistency of X-ray and optical data. This imply to use only the BeppoSAX data for the 2–10 keV band and optical magnitudes in the R band.

**3. Normalization of the X-ray data**

In order to be able to compare the flux in the 2–10 keV band, we normalized the data to a common distance corresponding to a redshift of 1. We applied the following corrections: distance-luminosity (relative to the target distance), normalization to the 2–10 keV band in flux, and time dilation of the temporal scale. These corrections have been computed using the spectral index provided for the GRB source spectra when available. When the

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**Fig. 1a and b.** X-ray light curves of the burst sample **a** no correction has been applied, and **b** with distance, flux, and time dilation normalization to a redshift of 1 applied

burst X-ray spectrum was not known with enough accuracy, we used the value of 1, which seems to be consistent with the majority of the burst afterglow spectra (Costa 1999). We took as a base a flat universe with  $\Omega_{\Lambda} = 0.7$ . Table 2 details the resulting correction factors applied to the data.

Fig. 1a displays the X-ray afterglow raw data. The same sample is displayed Fig. 1b, but we applied the corrections shown in Table 1 in order to plot the light curves for a standard distance corresponding to a redshift of 1.

#### 4. Optical data

The same corrections have been computed for the optical data in the R band from the same sources, of course taking into account the wavelength. At optical wavelengths the galactic absorption may be somewhat high. We used the work of Schlegel et al. (1998) to derive an additional correction to apply to the data. The resulting factors are shown on Table 3.

The raw optical light curves are displayed on Fig. 2a. On Fig. 2b we plotted the corrected light curves, again for a standard distance corresponding to a redshift of 1. As it can be seen,

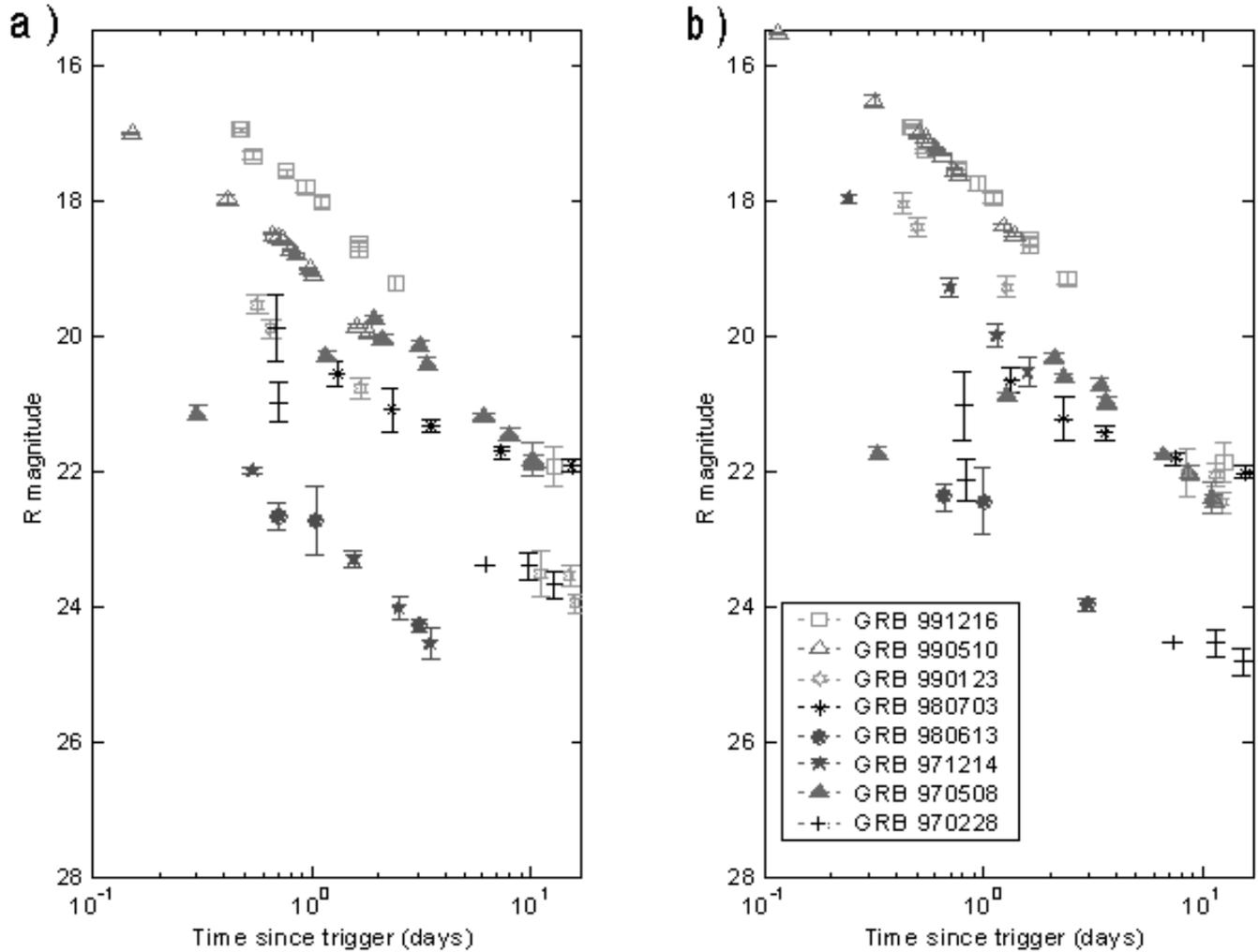
**Table 2.** X-ray correction factors

Source	Redshift	Correction factors	
		Distance	flux
GRB 970228	0.695	0.41	0.85
GRB 970508	0.835	0.64	0.92
GRB 971214	3.42	20.2	2.21
GRB 980613	1.096	1.25	1.05
GRB 980703	0.966	0.92	0.99
GRB 990123	1.61	3.23	1.19
GRB 990510	1.62	3.29	1.31
GRB 991216	1.02	1.05	1.01

the correlation observed at X-ray wavelengths vanishes almost completely, though the dispersion in magnitude is somewhat reduced.

#### 5. Discussion and conclusions

The X-ray data present, when corrected, an obvious correlation into two homogeneous groups. We find a mean slope of  $-1.6 \pm$



**Fig. 2a and b.** Optical light curves in the R band of the burst sample **a** no correction has been applied, and **b** with distance, flux, time dilation, and R absorption corrections applied, normalizing the distance to a redshift of 1

**Table 3.** Optical correction factors

Source	Correction factors		
	Distance	Flux	Absorption
GRB 970228	0.41	0.85	0.61
GRB 970508	0.64	0.92	0.05
GRB 971214	20.2	1.99	0.06
GRB 980613	1.10	1.25	0.22
GRB 980703	0.92	0.99	0.72
GRB 990123	3.23	1.22	0.04
GRB 990510	3.28	1.18	0.53
GRB 991216	1.05	0.09	1.67

0.2 for the most luminous subset, and  $-1.11 \pm 0.17$  for the less luminous afterglows.

If we try to interpret the observed correlation in the framework of the fireball model, we have to suppose that the differential density of the medium shocked by the fireball plays a secondary role. Let us refer to the Eqs. (7) to (12) of Sari et al. (1998). If we take  $p = 2.3 \pm 0.15$  for the index of the elec-

tron distribution power law, our results at X-ray wavelengths are compatible with the radiative case for the most luminous afterglow group, and with the adiabatic case for the less luminous group. This is also close to the standard value assumed by Sari et al. (1996) of  $p = 2.5$ .

However, the afterglow temporal evolution should be dependent on the surrounding medium density. It is clearly not the case here, though the burst sample is somewhat restricted. An explanation may be that the burst source surrounding medium has been “washed” before the shock. Of course, in the absence of any measure of the critical transition time of the afterglow light curve, it is impossible to conclude on the density parameter. However the correlation we present at X-ray wavelengths seems to be more compatible with a weak dependence on it, i.e. a somewhat low density medium. For the GRBs belonging to the most luminous group, this conclusion is consistent with the results obtained by Kumar (2000) for a sample of 7 bursts in the same 2–10 keV energy range. It is interesting to note that Kumar and Piran (2000) have shown that the width of the distribution function for the X-ray afterglow flux should be significantly

**Table 4.** R absorption for the optical data and comparison between the optical and X-ray power law decay indexes.

Source	R Absorption	Power law decay indexes	
		X-ray	Optical
GRB 970228	2.0	$1.28 \pm 0.11$	$1.10 \pm 0.1$
GRB 970508	-0.6	$1.02 \pm 0.21$	$1.17 \pm 0.04$
GRB 971214	2.0	$1.62 \pm 0.27$	$1.20 \pm 0.02$
GRB 980613	2.6	$1.05 \pm 0.37$	$1.00 \pm 0.01$
GRB 980703	0.0	$0.57 \pm 0.14$	$1.17 \pm 0.25$
GRB 990123	2.3	$1.44 \pm 0.20$	$1.44 \pm 0.07$
GRB 990510	0.5	$1.54 \pm 0.15$	$0.82 \pm 0.02$
GRB 991216	1.1	$1.78 \pm 0.01$	$1.22 \pm 0.04$

smaller than the spread in GRB fluences. The small number of sources we used in this work prevent from any firm conclusion, though GRB afterglows in the slow cooling regime may show a larger dispersion, which may be confirmed with the larger data set expected from HETE-2 and SWIFT.

It is difficult to find a firm explanation on the absence of correlation at optical wavelengths. Since the fireball model predicts that the ratio between the optical and X-ray luminosity should remain approximately constant during the early phases of the afterglow, we can try to explain the absence of correlation by an external reason, such as the absorption in the host galaxy of the GRB source. In this case the discussion of the above paragraph does not apply. Assuming that the X-ray and optical light curves have the same indexes, we computed the absorption coefficients tabulated in the second column of Table 4.

All values are relative to the magnitude of GRB 980703. For GRB 970508, we computed the correction on the decreasing part of the light curve after 1.5 day. This correction is negative for GRB 970508, the most luminous of our sample. Since the absorption occurs at UV wavelengths if we take into account the source redshift, these values are quite reasonable. They correspond to a rough column density of  $5 \times 10^{21}$  at most (Prehdel and Schmitt 1995). For these values of the host galaxy medium column density, the transmission coefficient of 2 keV photons is above 0.95 (Seward, 2000). In order to assess the role of the UV absorption, we computed the decay slopes at X-ray wavelengths, and we compare them with the rate of decay of the afterglow light in the R band. The results are given in the last two columns of Table 4.

The mean power law decay index at X-ray wavelengths is  $1.6 \pm 0.2$  for the most luminous afterglow group, and  $1.11 \pm 0.17$  for the less luminous one. At optical wavelengths the values are respectively  $1.17 \pm 0.13$ , and  $1.09 \pm 0.07$ . Though there is on average some difference at optical wavelengths also between the two groups, this feature cannot be considered as established.

We have presented in this letter a correlation between the afterglow light curves at X-ray wavelengths. While this behaviour is not observed in the optical domain, the decreasing rate may appear to be correlated in both energy ranges. The fireball model provides a framework to explain the X-ray light curves, while the host galactic absorption may shade the correlation in the visible range. The HETE-2 experiment, due to launch this year, will provide new data which may confirm this picture.

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