

# X-ray/optical analysis of the cluster of galaxies Abell 1300: indications of a post-merger at $z = 0.31$

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**Abstract.** We present ROSAT PSPC and HRI observations of the distant cluster A 1300 ( $z = 0.3071$ ) in conjunction with optical data, to investigate the physics of the intra-cluster medium and the dynamical state of the cluster. By means of a multi-resolution wavelet analysis, we find evidence for structures in the X-ray emission of this very luminous ( $L_X \sim 1.7 \times 10^{45}$  erg s<sup>-1</sup> in the 0.1–2.4 keV band) and massive ( $M_{\text{tot}} \sim 1.3 \times 10^{15} M_\odot$  at a radius  $\sim 2.2$  Mpc) cluster. This cluster is significant as we appear to be witnessing the end of a merger which occurred at an early epoch.

**Key words:** galaxies: clusters: general – galaxies: clusters: individual: Abell 1300 – cosmology: observations – dark matter – X-rays: galaxies

## 1. Introduction

A significant fraction of clusters of galaxies appears to be undergoing mergers: 30 % of those surveyed with EINSTEIN (Forman & Jones, 1990), although this is certainly an underestimate as deeper ROSAT exposures show evidence for mergers even in apparently relaxed clusters such as Coma (White et al., 1993) or A2256 (Briel et al., 1991). How this fraction evolves with redshift is still an open question, but the answer is crucial as it provides important cosmological constraints and has implications for hierarchical clustering.

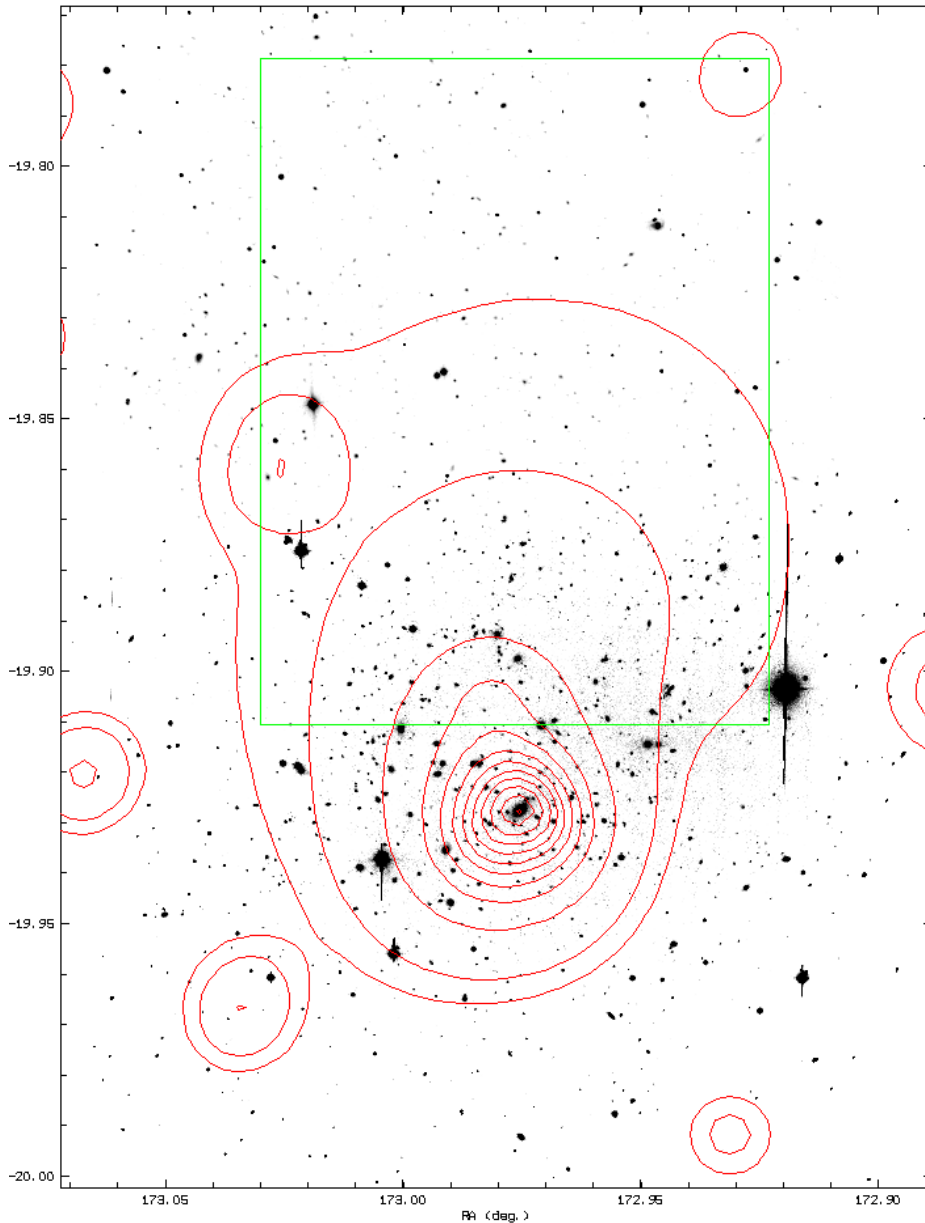
Some of the substructure we see in nearby clusters — such as Coma — implies only minor disturbances related to comparatively small mass components. The detection of substructure in distant clusters, however, naturally implies much larger disturbances because of the lower spatial resolution. If this could be quantified (and discriminated from other environmental effects) it could be used as a measure of the degree of cluster

growth as a function of time, and so constrain fundamental cosmological parameters, i.e., the amount of dark matter, index of the primordial fluctuation power spectrum and  $\Omega$ . However, as observational difficulties become greater with distance, multi-wavelength studies of particularly luminous clusters are required to be able to discern any evidence of a merger.

In a high resolution multi-wavelength campaign (IR, radio, optical and X-ray) initiated in order to study in detail the properties of distant bright X-ray clusters newly discovered in the ROSAT All-Sky Survey (RASS) (Pierre et al., 1994a), a sample of  $\sim 10$  clusters having X-ray luminosities greater than  $10^{44}$  erg s<sup>-1</sup> was selected, covering the redshift range  $0.1 < z < 0.3$  (Pierre et al., 1994b). The cluster of galaxies A 1300, which is described as a richness class 1 object (Abell et al., 1989), was found to be among the most luminous ( $L_X > 10^{45}$  erg s<sup>-1</sup> in the ROSAT hard band), most extended ( $\sim 3$  Mpc) and most distant ( $z = 0.3071$ ) in our sample.

Subsequent detailed spectroscopy and photometry of the main body of the cluster yielded a high velocity dispersion ( $\sigma_v \sim 1200$  km s<sup>-1</sup>; Pierre et al., 1997, hereafter Paper I). This large value is hard to reconcile with the picture of a single well relaxed cluster. Although the histogram of the velocity dispersion did not show significant substructure, the number density contours indicated the presence of several sub-groups of galaxies, especially East and North of the cluster center. These facts taken together suggest that A 1300 has undergone a merger at some stage. This hypothesis is now investigated in detail with additional optical data for the Northern part of the cluster combined with a study of the X-ray morphology. An associated radio study will follow (Reid et al., 1997, in preparation; hereafter Paper III).

The paper is structured as follows: the optical observations are presented in Sect. 2, the X-ray observations in Sect. 3, and the merging hypothesis is then discussed in Sect. 4 where a reference is made to the first results of the radio study.



**Fig. 1.** Optical field of A 1300, observed at the CFHT in R band for the southern part, and at the NTT in V band for the northern part. Overlaid are the ROSAT PSPC contours in the energy range 0.4–2.4 keV with levels [2.3, 5 (16) 165]  $\times 10^{-7}$  counts  $s^{-1}$  arcsec $^{-2}$ . The first contour corresponds to a  $3\sigma$  level detection. The X-ray image has been filtered using a multi-resolution wavelet analysis and pointing corrected to align the X-ray peak with the cD galaxy (see Sect. 3.3). The box shows the northern area covered by the new spectroscopic data (see Fig. 2). The displayed field is  $11'$  ( $\alpha$ )  $\times$   $15'$  ( $\delta$ ).

Throughout the paper we adopt a cosmological model with  $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$  and  $q_0 = 1/2$ . Celestial coordinates are in J2000.

## 2. Optical data

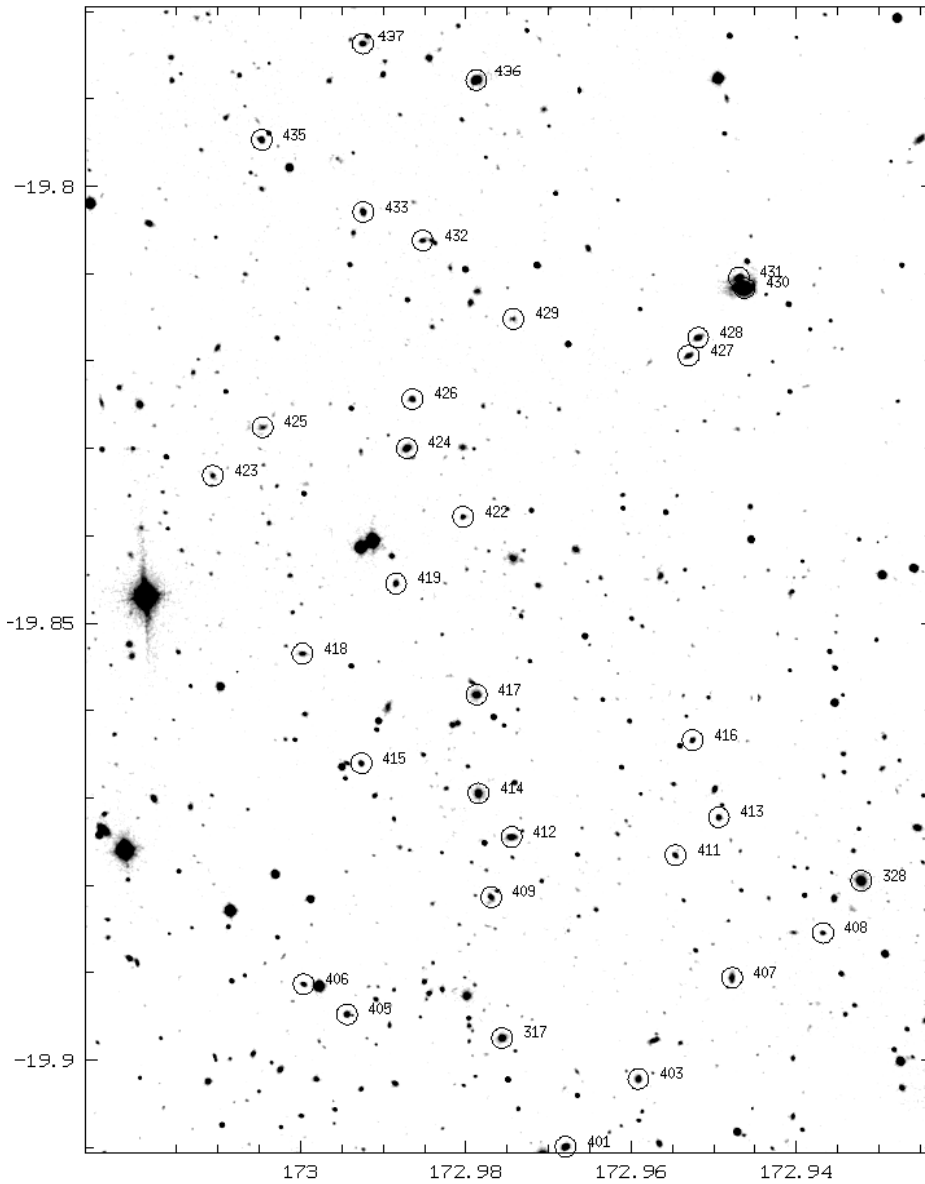
### 2.1. Observations

The data were collected during a three night run at the New Technology Telescope (NTT) at La Silla in February 1996, where four clusters were observed with the ESO Multi-Mode Instrument (EMMI). The selected mode was Red Imaging and Low Dispersion spectroscopy (RILD) and the CCD was the Tek #36 ( $2k \times 2k$  pixels of  $24 \mu\text{m}$ ). This is a thinned, back-illuminated CCD with a quantum efficiency of better than 60% between 400 and 800 nm. The observing configuration provides a pixel size of  $0''.27$  over a field of view of about  $9' \times 9'$ .

The seeing was good for the whole run with a stellar FWHM of  $\sim 1''.1$ . From our previous CFHT data (Paper I) there was no obvious fall-off in galaxy surface density towards the north, and the X-ray emission is also extended in this direction (Fig. 1), so we wanted to investigate this area in more detail. One mask of 37 spectra was taken for a field located some  $6'$  north of the cD galaxy (Fig. 2). Observations are summarized in Table 1.

### 2.2. Quick photometric analysis

In order to compare the X-ray map with the light distribution of the cluster, a photometric catalog of the galaxies was produced from the 5-minute V-band image of the present target area. The photometric analysis was performed by means of the SExtractor package (Bertin & Arnouts 1995) in the same way as Paper I, but adapted to our data. The V image was first cleaned of cosmic ray events (local clipping and median filtering), then slightly



**Fig. 2.** Finding chart for galaxies with new measured redshifts. North is up and east is to the left; the field of view is  $6'.2$  ( $\alpha$ )  $\times$   $7'.9$  ( $\delta$ ). This area covers the northern extent of the hard X-ray image (see box in Fig. 1). The reference numbers refer to Table 3, and to Paper I for objects 317 and 328.

smoothed by a  $\sigma = 1$  pixel Gaussian (comparable with the seeing), and finally the background was estimated using a  $32 \times 32$  pixel mesh. Source detections were claimed if at least 9 adjacent pixels were above a threshold corresponding to 3 times the local background level. The photometric calibration was done using the UBVRI Johnson-Kron-Cousins system (Landolt 1992) in SA98, with the stars observed during the same run. Estimates of the photometric errors were taken directly from the SExtractor analysis, and are summarized in Table 2.

The catalog is estimated to be complete to  $V = 23$ . The fraction of galactic stars was estimated using the Robin et al. (1995) model, and some 130 stars with  $10 \leq V \leq 23$  are expected in our field. This means that detected objects within this magnitude range having a SExtractor classification  $< 0.95$  may be assumed to be galaxies, i.e., 702 objects (see Fig. 3). Changing the threshold does not affect the outcome significantly because

most of the galaxies are well separated from stars (3/4 of the galaxies fall below 0.2).

Using all detected galaxies brighter than  $V = 23$ , we derive two isopleth maps, one showing galaxy number density (Fig. 4) and the other galaxy luminosity density (Fig. 5). These plots cannot be compared directly with those of Paper I (where only galaxies detected in both R and B were selected), but the overall morphology in the overlapping region is the same. We clearly see the northern extension which reaches up to  $5'$  from the center, i.e. 1.7 Mpc at the cluster redshift. Moreover, there seems to be a further galaxy concentration to the northeast (labelled 1 in Fig. 4).

### 2.3. Spectroscopy

Grism #3 was used for the spectroscopy. It has a zero deviation at  $4600 \text{ \AA}$ , covers approximately  $3850\text{--}8700 \text{ \AA}$ , and gives

**Table 1.** Observing summary for A 1300. The cD position (J2000) is 11h31m54.1s,  $-19^{\circ}55'40''$ . The 1993 run yielded 64 spectra (see Paper I) and the new run, shown in bold, 37.

	Field center	Integration time (min)	Observing run
<i>Photometry</i>			
R images	cD	15, 10	CFHT 93
B images	cD	2×25	CFHT 93
<b>V image</b>	<b>6' north of the cD</b>	<b>5</b>	<b>NTT 96</b>
<i>Spectroscopy</i>			
1st mask	cD	45, 40	CFHT 93
2nd mask	cD	75, 50	CFHT 93
3rd mask	3' north of the cD	60, 75	CFHT 93
<b>4th mask</b>	<b>6' north of the cD</b>	<b>2×55</b>	<b>NTT 96</b>

**Table 2.** Maximum and mean RMS errors in the photometric data, as estimated by the SExtractor program.

Magnitude	Mean RMS error	Max. RMS error
$V \leq 20$	0.011	0.05
$20 < V \leq 21$	0.018	0.065
$21 < V \leq 22$	0.033	0.17
$22 < V \leq 23$	0.065	0.29
$V > 23$	0.11	0.51

a dispersion of 2.3 Å/pixel. The slit has a width of  $1''.87$ , i.e. 6.9 pixels, yielding a resolution of  $\sim 16$  Å FWHM. Wavelength calibration was done using internal Helium and Argon lamps and subsequent reduction was performed as described in Paper I. Redshifts were measured by a cross-correlation method following Tonry and Davis (1979) and implemented in the MIDAS environment. Of the 37 spectra, two were chosen to be in common with those of the previous CFHT run to check internal consistency. Objects 317 and 328, which were previously found to have redshifts of  $z = 0.3076$  and  $z = 0.2575$  respectively (see Paper I), are in very good agreement ( $z = 0.3080$  and  $z = 0.2578$ ); the mean difference of  $\sim 80$  km s $^{-1}$  is well within the quoted spectroscopic uncertainties.

The results from the cross-correlation analysis for faint spectra were checked by hand and the presence of a few conspicuous absorption lines were required for them to be included in the final list which is presented in Table 3. Heliocentric correction has not been applied, but is negligible at this resolution.

Cluster membership was assessed using a  $3\sigma$  clipping method and 10 of the 37 galaxies were found to belong to the cluster (one of them in common with Paper I). The northernmost cluster galaxy (#424) is 5.9 (i.e.  $\sim 2$  Mpc at the cluster redshift) from the cD galaxy. The cluster redshift distribution is displayed in Fig. 6. As already noted in Paper I, while the

**Table 3.** Spectroscopic analysis of the northern part of A 1300:

Column 1: internal reference number to Fig. 2 (and Paper I for galaxies 317 and 328).

Column 2 & 3: RA and Dec (J2000), decimal degrees. Galaxy positions were determined from the V image and should have an accuracy of  $\sim 0''.7$  rms.

Column 4: redshift

(\*) signifies the presence of emission lines:

401: [O II], [O III], H $\beta$ , H $\alpha$ , [N II], [S II]

405: [O II], [O III], H $\gamma$ , H $\beta$

435: [O II], [O III], H $\beta$

(†): two interacting galaxies. No emission lines.

Column 5:  $\Delta z$  is the internal measurement error and is related to the correlation coefficient  $c$  by the formula  $\Delta z = k/(1+c)$  where  $k \sim 150$  km s $^{-1}$  was determined by the Tonry and Davis (1979) method. Comparison with redshifts for galaxies in common with Paper I suggests that  $\Delta z$  is conservative.

Column 6: redshift measurement quality:

1: highest peak in the correlation function,

2: interactive measurement using lines,

9: spectrum too noisy to be verified by hand.

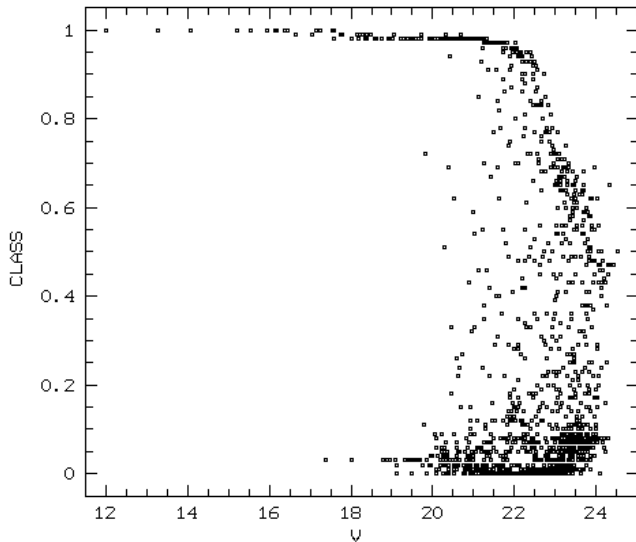
Column 7: V magnitude:

(1) indicates the presence of neighbors, bright and close enough to bias magnitude determination significantly ( $\Delta V \sim 0.1$ ).

(2) indicates that the object was originally blended with another one.

Column 8: Cluster member galaxy ( $3\sigma$  clipping method).

ID	RA(J2000)	Dec(J2000)	$z$	$\Delta z$	Q	V	member
401	172.9679	-19.9099	0.2413 <sup>(*)</sup>	0.0010	2	19.63	N
403	172.9591	-19.9022	0.0029	0.0011	9	20.28	N
317	172.9756	-19.8975	0.3080	0.0007	1	19.48 <sup>(2)</sup>	Y
405	172.9944	-19.8948	0.2955 <sup>(*)</sup>	0.0012	2	20.10	Y
406	172.9996	-19.8913	0.3098	0.0010	1	20.61	Y
407	172.9478	-19.8906	0.2572	0.0012	9	20.14	N
408	172.9368	-19.8855	0.3696	0.0012	9	21.24	N
409	172.9770	-19.8814	0.2710	0.0017	2	20.31 <sup>(1)</sup>	N
328	172.9322	-19.8795	0.2578	0.0007	1	18.95	N
411	172.9547	-19.8766	0.3079	0.0010	1	20.53	Y
412	172.9745	-19.8745	0.2548	0.0011	1	19.94 <sup>(1)</sup>	N
413	172.9494	-19.8722	0.2560	0.0009	1	20.64	N
414	172.9785	-19.8695	0.3054	0.0007	1	19.45	Y
415	172.9927	-19.8660	0.3075	0.0010	1	20.66	Y
416	172.9526	-19.8634	0.3111	0.0009	1	20.72	Y
417	172.9787	-19.8582	0.3013	0.0010	1	19.52 <sup>(1,2)</sup>	Y
418	172.9998	-19.8535	0.2773	0.0011	1	20.63	N
419	172.9885	-19.8455	0.1727	0.0009	9	20.44	N
422	172.9804	-19.8378	0.3070	0.0012	1	21.00	Y
423	173.0107	-19.8331	0.2158	0.0010	9	20.93 <sup>(1)</sup>	Y
424	172.9872	-19.8300	0.3074	0.0009	1	19.57 <sup>(2)</sup>	Y
425	173.0046	-19.8276	0.1851	0.0011	9	20.80 <sup>(1,2)</sup>	N
426	172.9865	-19.8244	0.2777	0.0010	9	20.58	N
427	172.9531	-19.8194	0.1606	0.0012	9	20.53	N
428	172.9519	-19.8173	0.0482	0.0011	9	20.27	N
429	172.9743	-19.8152	0.3878	0.0011	9	21.26	N
430	172.9463	-19.8117	0.1578 <sup>(†)</sup>	0.0006	1	17.38 <sup>(2)</sup>	N
431	172.9470	-19.8105	0.1592 <sup>(†)</sup>	0.0006	2	20.07 <sup>(2)</sup>	N
432	172.9853	-19.8062	0.3683	0.0010	1	20.85 <sup>(1,2)</sup>	N
433	172.9924	-19.8030	0.2311	0.0007	1	20.06	N
435	173.0047	-19.7947	0.1519 <sup>(*)</sup>	0.0009	2	20.26 <sup>(1)</sup>	N
436	172.9788	-19.7879	0.2311	0.0006	1	18.88	N
437	172.9925	-19.7837	0.2294	0.0007	1	20.27 <sup>(1,2)</sup>	N



**Fig. 3.** SExtractor object classification as a function of V magnitude for the A 1300 field shown in Fig. 2. The SExtractor star/galaxy classifier ranges from 0 (galaxy) to 1 (star). All objects up to class = 0.95 are expected to be galaxies, i.e. 702 out of 832 objects up to  $V = 23$ , the magnitude completeness (see text).

histogram is very broad, it does not show any significant substructure. Indeed, the whole dataset is well fitted by a Gaussian centered on  $z = 0.3071$  with  $\sigma = 0.0052 \pm 0.0005$ , giving a radial velocity dispersion of  $\sim 1200 \text{ km s}^{-1}$ .

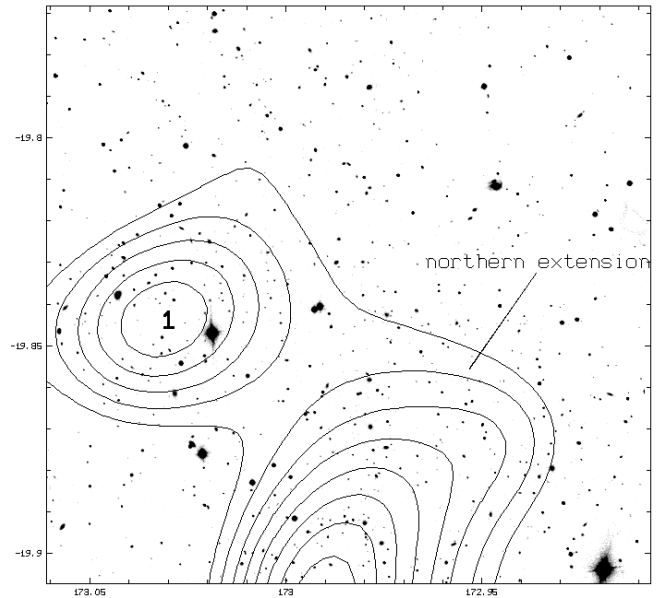
The velocity distribution was investigated in more detail using the ROSTAT package (Beers et al., 1990; Bird & Beers, 1993). Skewness and kurtosis ( $-0.092$  and  $2.683$  respectively) strongly suggest a Gaussian distribution (formally: 0 and 3 respectively). We have also investigated the presence of substructures in  $(\alpha, \delta, z)$  space by producing a velocity map. No obvious signal was detected although more redshifts are required for a proper statistical analysis.

### 3. X-ray data

#### 3.1. Observations

Pointed ROSAT Position Sensitive Proportional Counter (PSPC) observations were obtained in the period June 27–28 1993, with four exposures totalling 8572 sec. The PSPC has an on-axis spatial resolution of  $\sim 30''$  (FWHM), is sensitive to photons with energies between 0.1 and 2.4 keV, and has an energy resolution of about 40% at 1 keV (Briel et al., 1988). Data were analyzed using the EXSAS package in MIDAS.

In addition, a 11,000 s High Resolution Imager (HRI) observation of the central region of A 1300 was obtained on July 9 1994 (3,000 s) and on June 14–15 1995 (8,000 s). Exposure time was inadequate to reach the desired signal-to-noise ratio but was sufficient to show an elongated structure which peaks on the cD galaxy and extends to the south east. The contour image is shown in Fig. 7 overlaid on the R-band CCD image.



**Fig. 4.** Galaxy isodensity contours for galaxies with  $V < 23$  overlaid on the CCD image of the field centered  $6'$  N of the cD galaxy (Fig. 2). A multi-resolution wavelet analysis was applied. The first contour corresponds to  $10 \text{ galaxies arcmin}^{-2}$  ( $3\sigma$  level above background), and the contour increment is  $1.2 \text{ galaxies arcmin}^{-2}$ . A northern extension is clearly seen in the galaxy distribution, corresponding to the northern X-ray extension. Moreover, a further galaxy concentration, labelled 1, is clearly detected (see text). The image size is  $8.8 (\alpha) \times 8.5 (\delta)$ .

The following sections will concentrate on the analysis of the PSPC image.

#### 3.2. Source detection

A maximum likelihood search for discrete sources was carried out over the PSPC 0.1–2.4 keV image and the results are listed in Table 4. The X-ray point source positions were correlated with the SIMBAD<sup>1</sup> and NED<sup>2</sup> on-line catalogs but no identifications were found. Comparison with the COSMOS<sup>3</sup> source list, derived from a circular region of radius  $20''$  centered on each object from Table 4, produced a number of possible optical identifications. These are listed in Table 5.

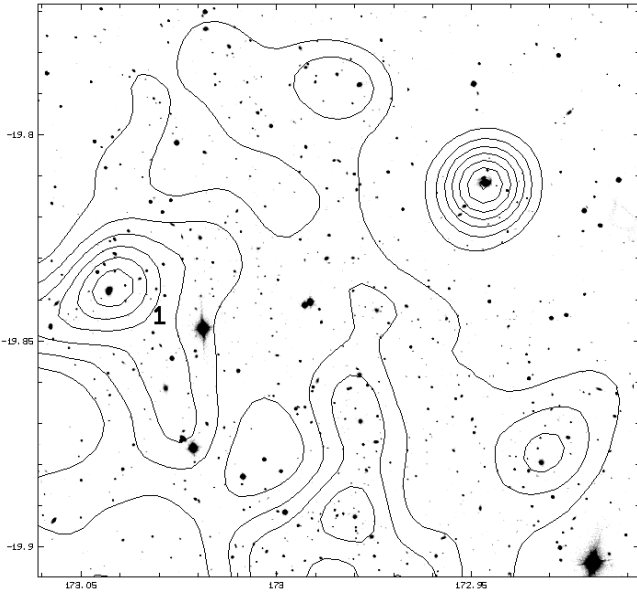
#### 3.3. X-ray morphology

A fundamental aspect of morphological studies of the intracluster medium (ICM) is the search for substructure. Because of its shorter relaxation time, the X-ray gas is expected to trace the shape of the gravitational potential more closely than galaxies. The X-ray morphology, therefore, is better suited to the investigation of substructures and merging events as well as providing insights into the dark matter distribution.

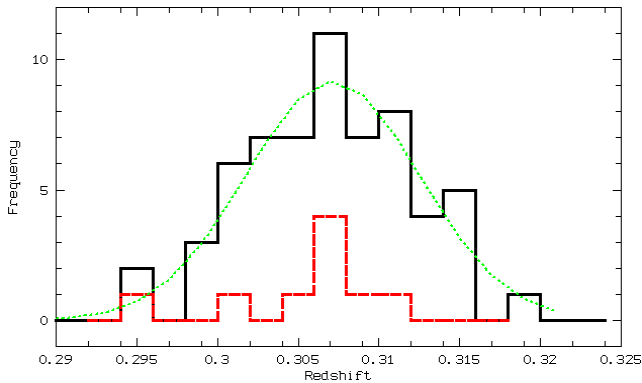
<sup>1</sup> <http://simbad.u-strasbg.fr/Simbad>

<sup>2</sup> [telnet://ned@ned.ipac.caltech.edu/](mailto:telnet://ned@ned.ipac.caltech.edu/)

<sup>3</sup> [http://xweb.nrl.navy.mil/www\\_rsearch/RS\\_form.html](http://xweb.nrl.navy.mil/www_rsearch/RS_form.html)

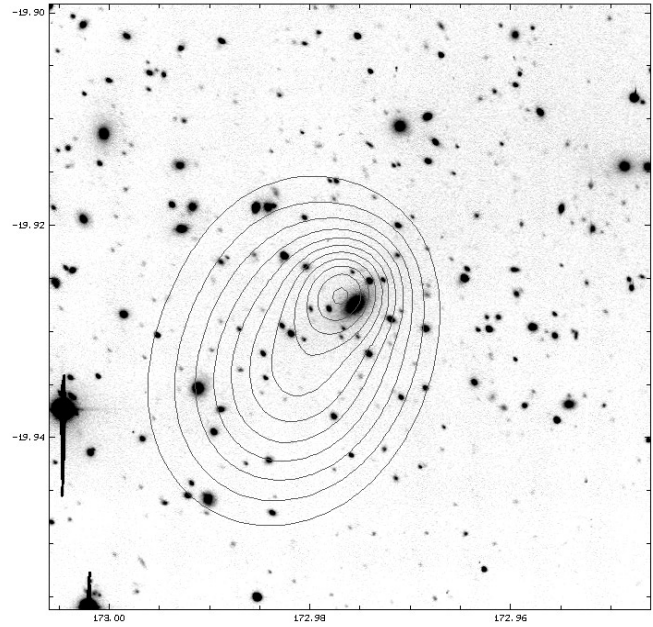


**Fig. 5.** Luminosity weighted density contours for the data of Fig. 4. The filtering function has a  $\sigma$ -value of  $27''$ . The contour increment is constant in flux, with the first and last contours corresponding respectively to 26.5 and 25.0 mag arcmin $^{-2}$ . The principal extended structure is  $5'$  away from the cluster center, i.e. 1.7 Mpc. The label 1 indicates the position of a clump detected in the galaxy density image; it is contaminated here by a bright galaxy. The very bright point-like structure to the northwest arises from two interacting galaxies which do not belong to the cluster (see Table 3). Image size is  $8'.8$  ( $\alpha$ )  $\times$   $8'.5$  ( $\delta$ ).



**Fig. 6.** Cluster galaxy redshift histogram (bin size  $\Delta z = 0.002$ ); the mean redshift is 0.3071 and the velocity dispersion  $\sim 1200$  km s $^{-1}$ . The solid line represents the whole dataset (Paper I + present data) while the dashed line shows the nine new spectra. Dotted line is the best Gaussian fit.

Because of the very high temperature of the ICM ( $10^7$ – $10^8$  K, Sarazin 1986), the energy is mostly emitted in the range 1–10 keV. Moreover, the soft ROSAT band, i.e.  $\leq 0.4$  keV, is strongly contaminated by background emission, predominantly from the Galaxy. Consequently, the photon file was binned into separate  $5''$  pixel images, with energies in the range 0.4–2.4 keV (expected to belong to the cluster) and 0.1–0.4 keV. These im-



**Fig. 7.** ROSAT HRI image of the central region of A 1300 ( $t_{\text{exp}} = 10^4$  s,  $\sim 200$  excess photons above background) overlaid on the CCD image. The contours are equally spaced from  $3$ – $8 \times 10^{-6}$  counts s $^{-1}$  arcsec $^{-2}$ . The offset between the X-ray peak and the cD galaxy of  $\sim 5''$  is assumed to be due to satellite attitude error. Image size is  $3'.6 \times 3'.4$ .

ages were then filtered using a multi-resolution wavelet analysis program (Starck et al., 1996). This method suppresses photon noise and restores structures on different scales, in this case with a confidence level of 99.9%. Fig. 1 shows the hard X-ray contours overlaid on the optical image of the cluster.

Alignment of the optical images was obtained by computing the transformation parameters between CCD coordinates and celestial coordinates using ten stars with equatorial coordinates measured on the COSMOS UKST Southern Sky Catalog. The standard deviation of the residuals was less than  $0''.7$ . Because of the small number of photons in all of the point-like X-ray sources that were detected (about twenty for the brightest) we were unable to check the attitude solution of the X-ray image. We therefore decided to align the peak in the hard band X-ray image with the optical cD galaxy. This corresponds to a shift in the X-ray center of  $9''$  west and  $9''$  north. This choice appears to be compatible with the location of some of the point sources in the field with faint counterparts in the digitized sky survey<sup>4</sup>. Further support for the above registration came after multi-resolution wavelet filtering which showed that the very soft source (number 8 in Table 4 and the best spatially determined) now coincides with the stellar object (possibly a QSO) identified in Table 5.

Contours of the filtered soft- and hard-band images of A 1300 are overlaid in Fig. 8. The hard-band image is clearly asymmetric and has a higher peak count rate. There is also a significant displacement of  $\sim 23''$  between the soft and hard

<sup>4</sup> <http://skview.gsfc.nasa.gov/skyview.html>

**Table 4.** Discrete X-ray sources in the 45' diameter field of A 1300 detected with a maximum likelihood greater than 10. Coordinates refer to the normal attitude solution, which is likely to have a 9'' east and 9'' south shift with respect to absolute celestial coordinate (see text).

ID	RA(J2000) h m s	Dec(J2000) ° ' "	$f_x$ counts s <sup>-1</sup>	error
1	11 31 14.8	-19 45 17	0.0027	0.0009
2	11 31 48.7	-19 50 43	0.0026	0.0008
3	11 32 06.3	-19 51 35	0.0047	0.0010
4	11 31 52.5	-19 52 01	0.0042	0.0010
5	11 32 33.9	-19 52 47	0.0031	0.0008
6	11 31 58.3	-19 52 57	0.0121	0.0016
7	11 31 45.0	-19 53 48	0.0081	0.0014
8	11 31 54.6	-19 54 17	0.0410	0.0025
9	11 31 48.2	-19 54 51	0.0089	0.0014
10	11 31 12.2	-19 55 08	0.0042	0.0010
11	11 32 16.3	-19 55 20	0.0050	0.0010
12	11 31 54.8	-19 56 02	0.0896	0.0035
13	11 31 47.1	-19 56 11	0.0048	0.0011
14	11 32 08.4	-19 58 07	0.0043	0.0009
15	11 31 28.5	-19 59 00	0.0031	0.0008
16	11 30 54.7	-20 12 10	0.0056	0.0014

maxima. The soft source mentioned above (no. 8 in Table 4) can be seen immediately north of the hard X-ray peak. Interestingly, we detect an additional maximum in the hard band to the north-east. The multi-resolution analysis shows that this emission region (labelled “A” in Fig. 8) is not point-like and is slightly elongated north-south. It falls close in projection to the position of clump 1, detected in the galaxy isodensity image (Fig. 4); the offset of 1'7 corresponds to  $\sim 0.5$  Mpc at the cluster redshift.

On larger scales, the cluster shows a strong elongation to the north, with extended emission being detected as far as 8'3 north of the the peak ( $\sim 2.8$  Mpc at this redshift). In contrast the western extent is only 1'7 (0.5 Mpc). This important difference is displayed in Fig. 9 where the X-ray surface brightness profiles in three different directions are compared.

### 3.4. Surface brightness

Despite its clearly asymmetrical morphology, we did not have enough photons to further characterize the different sub-clumps. It is, however, possible to calculate a circularly averaged X-ray surface brightness profile of A 1300 from the hard band. We removed areas corresponding to the northern extension and all point-like sources from the hard image (restricted to 0.5–2.0 keV) and further assumed circular symmetry. We noticed that the profile fit was sensitive to a shift in the X-ray center, so we smoothed the data by binning into 30'' pixels. The profile was then fitted with an isothermal “ $\beta$ -model” (King, 1962; Cavaliere & Fusco-Femiano, 1976, 1981) of the form

$$S(R) = S_0[1 + (R/R_c)^2]^{0.5-3\beta} + \text{background}, \quad (1)$$

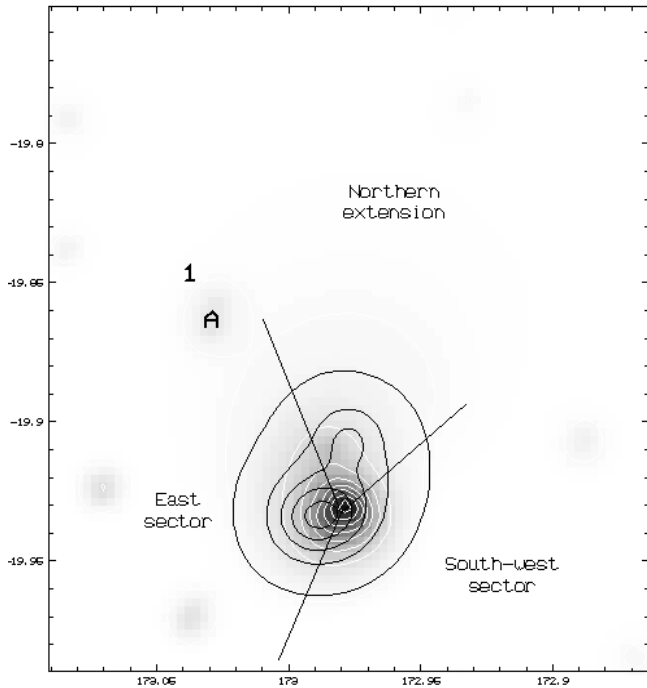
**Table 5.** Possible COSMOS optical identifications for the discrete X-ray sources detected in the ROSAT field. As some sources fall outside the CCD field, we adopt the COSMOS magnitude for uniformity. Column 1 refers to the X-ray ID, columns 2 & 3 give the optical position of potential optical counterparts within a 20'' radius, column 4 gives the ROSAT-optical offset in arcsec, column 5 gives the COSMOS  $b_J$  magnitude, and column 6 gives the COSMOS classification: s = star, g = galaxy, f = faint source.

ID	RA(J2000) h m s	Dec(J2000) ° ' "	$\Delta r$ "	$b_J$	class
1	11 31 14.28	-19 45 19.0	7.8	20.0	s
	11 31 16.07	-19 45 15.4	18.0	16.7	s
2	11 31 49.54	-19 50 39.8	12.6	22.2	f
3	11 32 06.81	-19 51 41.2	9.6	20.5	g
	11 32 06.45	-19 51 15.2	19.8	17.7	s
4	11 31 53.79	-19 52 05.7	18.6	21.5	s
5	11 32 34.61	-19 52 45.3	10.2	18.0	s
6	11 31 59.72	-19 52 53.8	20.4	18.1	s
7	11 31 45.69	-19 53 59.3	15.0	22.6	f
	11 31 43.89	-19 53 36.5	19.2	23.0	f
8	11 31 53.98	-19 54 07.7	13.0	20.4	s
9	11 31 47.38	-19 54 52.8	12.0	19.3	s
	11 31 47.58	-19 54 52.7	9.0	20.6	g
	11 31 47.00	-19 54 53.1	16.8	21.3	g
11	11 32 16.16	-19 55 20.9	2.4	20.0	s
	11 32 15.33	-19 55 15.5	14.4	20.1	g
12	11 31 53.69	-19 55 56.7	16.2	22.2	f
	11 31 55.63	-19 55 50.2	16.8	22.8	f
	11 31 54.70	-19 55 41.8	20.4	22.2	f
14	11 32 08.78	-19 58 09.7	6.0	22.2	f
15	11 31 28.31	-19 59 02.9	4.2	19.0	s

simultaneously deconvolving by the PSPC PSF at 1 keV (Hasinger et al., 1992).

The fit, convolved with the PSPC resolution, is overlaid on the data in Fig. 10. The values of the fitted parameters were:  $S_0 = (1.44 \pm 0.2) \times 10^{-5}$  counts s<sup>-1</sup> arcsec<sup>-2</sup> (at a 95% confidence level),  $\beta = 0.64 \pm 0.03$ ,  $R_c = 43 \pm 5''$  ( $\sim 240$  kpc at the cluster redshift; both errors given at a 99% confidence level), and a background value of  $8.1 \times 10^{-8}$  counts s<sup>-1</sup> arcsec<sup>-2</sup>. Note that  $\beta$  and the core radius are only representative of the circularly averaged surface brightness. Their values are very close to average values for a large number of clusters studied by Jones & Forman (1984). The fact that  $\beta$  is less than unity indicates that the energy per unit mass in the gas is higher than in the galaxies and that the gas density falls off less rapidly with radius than the galaxy density. However, as Fig. 9 shows, there are significant differences in the profiles in different directions.

If we continue to assume circular symmetry and use the averaged characteristics for the cluster, the surface brightness (Eq. 1) can be deprojected analytically. Furthermore, the observed surface brightness is weakly dependent on the gas temperature (this is justified because of the expected high temperature of this very bright cluster and the poor temperature sensitivity of



**Fig. 8.** ROSAT PSPC images of A 1300. The soft image (0.1–0.4 keV) shown in equally spaced black solid contours from  $0.35\text{--}1.75 \times 10^{-7}$  counts  $\text{s}^{-1} \text{arcsec}^{-2}$ , is overlaid on the hard image (0.4–2.4 keV) shown in greyscale with equally spaced white contours for  $0.47\text{--}16.3 \times 10^{-6}$  counts  $\text{s}^{-1} \text{arcsec}^{-2}$ . Filtered images were obtained by a multi-resolution wavelet analysis (see text). Label “1” marks the position of a local clump in the galaxy isodensity image (Fig. 4), while “A” indicates a second luminosity maximum in the hard image. In each band the first contour level is  $3\sigma$  above the background level. The south-west and east sectors refer to the spectral analysis. Image size is  $13' \times 14'5$ .

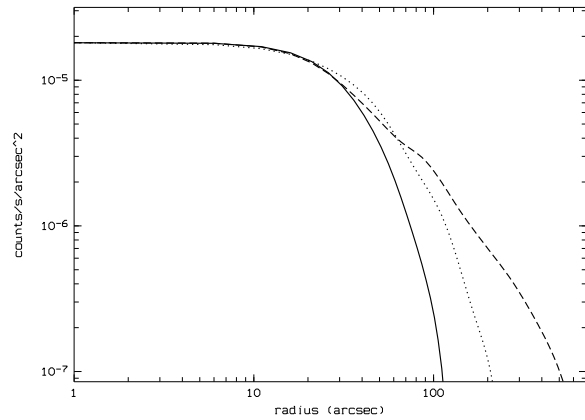
the ROSAT PSPC above 5 keV). In this case, the proton density distribution  $n_p(r)$  is given by

$$n_p(r) = n_p(0)[1 + (r/R_c)^2]^{-\frac{3\beta}{2}}. \quad (2)$$

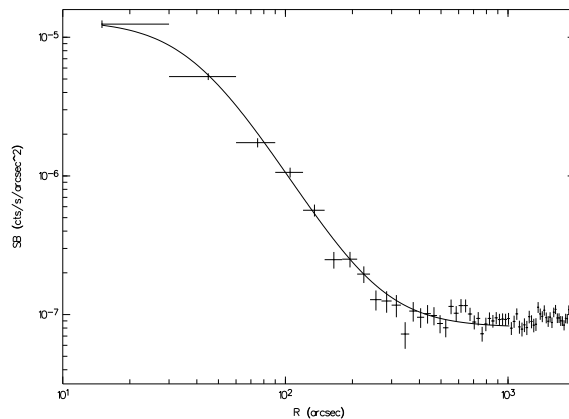
Using the above values for  $\beta$  and  $R_c$ , and normalizing the density profile according to the observed count-rates, we obtain  $n_p(0) = 8.5 \times 10^{-3} \text{ cm}^{-3}$ . This density will be adopted hereafter.

### 3.5. Spectral analysis

The ROSAT PSPC spectral range (0.1–2.4 keV) is not well suited to the temperature measurement of bright clusters. However, we fitted data between 0.1 and 2.4 keV with a Raymond & Smith (1977) thermal plasma model. The spectrum was binned such that each bin had a signal-to-noise ratio of 7, and the bins extended out to a radius of  $150''$  (850 kpc) from the center; a sector containing the soft point-like source was excluded. We fixed the metallicity at  $Z = 0.3$  (variations in the abundance did not affect the results by more than 10%) and obtained a Galactic neutral hydrogen column density of  $N_H = (4.1 \pm 0.9) \times 10^{20} \text{ cm}^{-2}$ ,



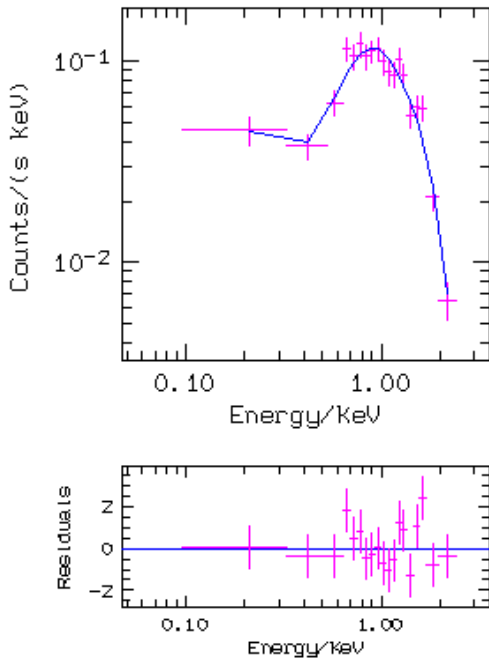
**Fig. 9.** The X-ray surface brightness asymmetry: profiles to the east (dotted), north (dashed) and west (solid) of the peak of the filtered hard image. The northern profile is clearly more elongated and falls off more gently than the east and west profiles. The steep fall off at low intensity is due to the wavelet filtering process. The background was estimated to be  $8.5 \times 10^{-8}$  counts  $\text{s}^{-1} \text{arcsec}^{-2}$ , so profiles were cut at this value. All structures are detected with a 99.9% confidence level.



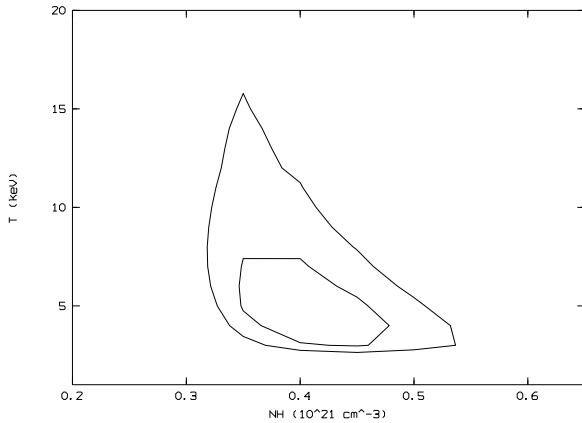
**Fig. 10.** X-ray surface brightness profile of A 1300 (circularly averaged) in the energy range 0.5–2.0 keV (channels 52–201). A bin size of  $30''$  was used to obtain a good signal-to-noise ratio. The solid line is the best fit to a King profile (Eq. 1) convolved with the PSF and overlaid on the data.

in good agreement with the value from Dickey & Lockman (1990) of  $N_H = 4.5 \times 10^{20} \text{ cm}^{-2}$ . With these constraints we obtained a temperature for the cluster of  $5 \pm 3 \text{ keV}$  ( $\chi^2 = 18$  with 16 variables). The fit is displayed on the spectrum in Fig. 11 and as  $\chi^2$  contours in Fig. 12).

The unusual morphology of A 1300 described above suggested that it may be worthwhile looking for spatially separated temperature components. We considered (i) the south-west sector (pa  $160^\circ\text{--}315^\circ$ , and (ii) the east sector (pa  $20^\circ\text{--}160^\circ$ ) (cf. Figs. 8, 9). Despite the small number of photons ( $\sim 500$  and  $\sim 600$  respectively), we find some indication of a temperature gradient. As before, spectra were binned so that each bin had a signal-to-noise ratio of 7, and fits were renormalized between



**Fig. 11.** Best fit of a Raymond & Smith thermal plasma model. The fit indicates a temperature of  $\sim 5 \pm 3$  keV.



**Fig. 12.** Contours of  $\chi^2$  as a function of column density  $N_H$  and temperature  $T$  for confidence levels of 68.3% and 95.4%

0.9 and 1.0 keV for comparison. In both cases, we found a Galactic  $N_H$  in good agreement with the previous value. The best fit for the south-western sector indicates a temperature of  $3 \pm 2$  keV, whereas that for the eastern sector suggests a much higher value ( $>7$  keV) (both temperatures at a 68% confidence level), but with very large uncertainties as expected with the ROSAT sensitivity. The temperature deduced from the galaxy velocity dispersion correlation ( $T_X = 10^{-3.22} \sigma^{1.35}$ ; Edge & Stewart, 1991b) is 9 keV, marginally consistent with our global determination but beyond the ROSAT spectral range. We tentatively adopt this value in the following sections.

### 3.6. Total luminosity

The total cluster luminosity was calculated from the fitted surface brightness profile. We integrated the ROSAT counts out to a radius  $r_{\text{cut}}$  of  $400''$  (2.2 Mpc) set by the detection limit of the circularly averaged surface brightness profile (see Fig. 10). This gives an overall count rate  $C$ , where

$$C = \frac{\pi S_0 R_c^2}{3/2 - 3\beta} \left[ (1 + r_{\text{cut}}^2)^{3/2 - 3\beta} - 1 \right], \quad (3)$$

of 0.25 counts  $\text{s}^{-1}$ . This count rate was converted into luminosity using a Raymond & Smith (1977) thermal model folded with the PSPC response provided by the EXSAS package, and assuming a thermal spectrum with  $T = 9$  keV, metallicity of 0.3 and a neutral hydrogen column density of  $4.5 \times 10^{20} \text{ cm}^{-2}$ . We obtain  $L_X = (1.7 \pm 0.1) \times 10^{45} \text{ erg s}^{-1}$  (at a 95% confidence level) in the 0.1–2.4 keV band. Using the empirical temperature-luminosity relation (Henry & Arnaud 1991, Edge & Stewart 1991a), this would indeed correspond to a temperature  $\sim 9$  keV, consistent with the estimate obtained from the galaxy velocity dispersion correlation (Sect. 3.5).

Although no excess X-ray emission is observed at the cluster center, we investigated the presence of a cooling flow. An estimate of the cooling time (Sarazin 1986) is given by

$$T_{\text{cool}}(\text{yr}) = 8.5 \times 10^{10} \left( \frac{T}{10^8 \text{ K}} \right)^{1/2} \left( \frac{n_p}{10^{-3} \text{ cm}^{-3}} \right)^{-1}. \quad (4)$$

Using the density profile determined above and  $T = 9$  keV, we derive  $R_{\text{cool}}$  which defines the region where the cooling time is smaller than the age of the universe at the cluster redshift. Depending on the value of the Hubble parameter  $h = H_0/100$ ,  $R_{\text{cool}}$  lies in the range 0–27'' for  $h = 0.75$ –0.5. For  $h > 0.75$ , the cooling time is greater than the age of the universe. Calculating the bolometric cluster luminosity ( $L_{\text{cool}}$ ) enclosed within  $R_{\text{cool}}$ , and assuming steady state isobaric gas cooling from 9 keV, we derive a mass flow (following Fabian et al., 1991) of

$$\dot{M}_{\text{cool}} = \frac{2}{5} \frac{\mu m_p}{kT} L_{\text{cool}} \sim 880 \left( \frac{L_{\text{cool}}}{10^{45}} \right) \left( \frac{T}{5 \text{ keV}} \right)^{-1} M_{\odot}/\text{yr} \quad (5)$$

This yields  $\dot{M}_{\text{cool}} < 400 M_{\odot}/\text{yr}$  for  $h > 0.5$ , but uncertainties are large (up to a factor  $\sim 2$ ).

### 3.7. Mass determination

We can obtain a robust estimate of the cluster mass by removing obvious regions of disturbance and assuming hydrostatic equilibrium for the rest of the cluster following Schindler (1996). We have a pressure gradient

$$\frac{dP}{dr} = -n\mu m_p \frac{d\phi}{dr}, \quad (6)$$

where  $P$  is given by the ideal gas law,  $P = nkT$ .

The mass integrated out to radius  $r$  is then

$$M(r) = -\frac{r^2 kT}{G\mu m_p} \left( \frac{1}{\rho_g} \frac{d\rho_g}{dr} + \frac{1}{T} \frac{dT}{dr} \right) \quad (7)$$

where  $\rho_g$  is the gas density. Using the density profile (Eq. 2) derived above, this yields

$$M(r) = -\frac{r^2 k}{G\mu m_p} \left( -\frac{3\beta r T}{r^2 + R_c^2} + \frac{dT}{dr} \right). \quad (8)$$

Since we expect the gas density profile to be accurate, the main error in the mass determination comes from the large uncertainty in the temperature. If we ignore any radial temperature dependence, we obtain a cluster mass of

$$M(r) = \frac{3\beta r^3 k T}{G\mu m_p (r^2 + R_c^2)} \quad (9)$$

which diverges with radius. As above, we integrate the mass out to a radius  $r_{\text{cut}}$  of 2.2 Mpc and obtain a total mass of  $M_{\text{tot}} \sim 1.3 \times 10^{15} M_{\odot}$ . Equation (2) gives a corresponding  $M_{\text{gas}}/M_{\text{tot}}$  ratio of  $\sim 30\%$ , in good agreement with previous determinations for rich clusters (Briel et al., 1992, Henry et al., 1993).

#### 4. Discussion and conclusions

For the first time in such a distant and rich cluster, high resolution multi-wavelength observations are enabling us to gain a deeper insight into the cluster dynamics.

Previous optical spectroscopy of A 1300 (Paper I) showed a very broad galaxy velocity dispersion ( $\sigma_v = 1200 \text{ km s}^{-1}$ ), hardly compatible with the picture of a relaxed cluster. In this paper, ROSAT PSPC observations, in conjunction with a multi-resolution wavelet analysis, have revealed an irregular morphology which is very extended to the north but falls off steeply to the south-west. Indications of non-isothermality following the brightness morphology were also found. The estimated total cluster mass is very high ( $M_{\text{tot}} \sim 1.3 \times 10^{15} M_{\odot}$ ) out to a radius of 2.2 Mpc, and the gas mass represents  $\sim 30\%$  of the total mass.

A 1300 is probably not in an equilibrium state as it shows signatures of a merger in the optical, X-ray and radio<sup>5</sup>. The properties of the cluster are summarised below.

- The overall X-ray emission is very extended (detected up to  $\sim 2.8$  Mpc from the cluster center in the north) and is clearly not spherically symmetric.
- Both the number density and luminosity-weighted galaxy counts show good correlation with the principal X-ray extension and substructure (clump 1), although the latter may not necessarily be at the same redshift as the cluster.
- The galaxy velocity dispersion is unusually high (consistent with an ICM temperature of  $\sim 9$  keV), and we found some indication of a temperature gradient in the PSPC data. The hotter regions, to the north and east of the cD galaxy, are also those of higher galaxy density, strengthening the reality of the temperature gradient.

- The X-ray luminosity is very high ( $\sim 1.7 \times 10^{45} \text{ erg s}^{-1}$  in the 0.1–2.4 keV band), supporting a high ICM temperature  $\sim 9$  keV on the basis of the empirical  $L_X:T$  relation (Henry & Arnaud, 1991).
- The present X-ray data exclude the presence of a strong cooling flow. However, Henriksen (1993) has found that cooling flows could form in rich clusters after a merger but before the galaxies have reached equilibrium.
- A short HRI exposure of the cluster center (Fig. 7) shows a clear elongated structure near the cluster core, peaking close to the position of the cD galaxy.
- Despite its high X-ray luminosity—and correspondingly large mass—we did not detect gravitational lensing effects (in the form of giant arcs or arclets) in the CFHT CCD images. This would suggest that the mass (luminous and dark) is not centrally peaked (as usually found in lensing clusters) but rather flat, as implied by the large X-ray core radius ( $\sim 230$  kpc), or possibly clumpy.
- A halo radio source has been discovered close to the center of the cluster and located in the same direction as the inner extension of the X-ray emission (Paper III).

At all wavelengths, therefore, it seems as though the results are consistent and confirm the exceptional nature of A 1300. Many of these properties may be compared with those found in more easily studied nearby clusters such as A754 ( $z = 0.0541$ : Henry & Briel 1995; Zabludoff & Zaritsky 1995) and A2256 ( $z = 0.0601$ : Briel & Henry, 1994), which also show substructure and are clearly far from hydrostatic equilibrium. Signatures of mergers and substructures have recently been seen as well in ROSAT X-ray data for some distant clusters, eg. Cl0016+16 (Neumann & Böhringer, 1997) and Cl0939+472 (Schindler & Wambsganss, 1996). These signatures were predicted in N-body/hydrodynamical simulations of hierarchical cluster evolution (Evrard, 1990a,b; Schindler & Muller, 1993; Roettiger et al., 1993) where clusters form through the merger of subclusters. Those simulations show that when subclusters collide, four effects can be observed over a time interval  $\sim 1h^{-1}$  Gyr : (i) the X-ray emission and the galaxy surface density do not coincide; (ii) the inner region of the ICM is compressed in the direction of the collision axis and this can appear as a perpendicular elongation in the X-ray surface brightness distribution; (iii) an anomalously hot component may be created in the ICM; and (iv) on the large scale an elongated shape with its axis parallel to the merger axis is predicted.

In the case of A 1300, the first two phenomena predicted by simulations are absent but we do detect a hot component and an elongated shape. The high galaxy velocity dispersion, together with the absence of clear substructure in the velocity histogram, suggests that this cluster has undergone a merger, but the merging phase may be nearly over. In this respect, A 1300 resembles the cluster A2142 ( $z = 0.09$ ) which displays a similar galaxy velocity histogram and velocity dispersion (Oegerle et al., 1995), integrated gas temperature (from GINGA data, White et al., 1994), and elliptical X-ray morphology and luminosity (Henry & Briel, 1996). Moreover, a detailed temperature map

<sup>5</sup> Observations of A 1300 by the MOST at 843 MHz and at higher resolution with the AT Compact Array at 1.4, 2.4, 4.8 and 8.6 GHz are presented in Paper III.

of A2142 (Henry & Briel, 1996) provides firm evidence for a non-relaxed ICM.

A 1300 is significant as we appear to be witnessing the end of a merger at a comparatively early epoch ( $z = 0.3$  corresponds to  $2/3$  of the age of the universe). An obvious question to ask is: do all clusters form at the same epoch or is cluster formation a continuous process? Comparing A 1300 with a similar nearby cluster such as A2142 argues strongly for the latter hypothesis. However, an extension of this work to much higher redshifts will only be possible with the new X-ray missions such as AXAF to provide the detailed morphology and XMM to provide spectral information. Radio observations (Paper III) reinforce the case for a recent merger in A 1300 and highlight the need for more sensitive high resolution X-ray observations to help understand the detailed relationship between radio morphologies and the cluster environment. Forthcoming ASCA data will provide a much better constraint on the ICM temperature and any gradients. Finally, a weak shear analysis on deeper optical images would be an ideal complement for mapping the large scale mass distribution and understanding its dynamical state.

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