

Optical-IUE observations of the gamma-ray loud BL Lacertae object S5 0716+714: data and interpretation^{*}

G. Ghisellini¹, M. Villata², C.M. Raiteri², S. Bosio³, G. De Francesco², G. Latini³, M. Maesano⁴, E. Massaro⁴, F. Montagni⁴, R. Nesci⁴, G. Tosti⁵, M. Fiorucci⁵, E. Pian⁶, L. Maraschi⁷, A. Treves⁸, A. Comastri⁹, and M. Mignoli⁹

¹ Osservatorio Astronomico di Brera, via Bianchi, 46, I-22055 Merate, Italy

² Osservatorio Astronomico di Torino, I-10025 Pino Torinese, Italy

³ Università di Torino, via Giuria 1, Torino, Italy

⁴ Istituto Astronomico, Università di Roma, Unità GIFCO-CNR Roma I, via Lancisi 29, Roma, Italy

⁵ Osservatorio Astronomico, Università di Perugia, Perugia, Italy

⁶ Space Telescope Science Institute, Baltimore, USA

⁷ Osservatorio di Brera, via Brera, 28, Milano, Italy

⁸ International School for Advanced Studies, Trieste, Italy

⁹ Osservatorio Astronomico di Bologna, Italy

Received 31 October 1996 / Accepted 3 June 1997

Abstract. We monitored the BL Lac object S5 0716+714 in the optical band during the period November 1994 – April 1995, which includes the time of a γ -ray observation by the Energetic Gamma Ray Experiment Telescope (EGRET) on February 14-28, 1995. The light curves in the R and B bands show fast fluctuations superimposed on longer timescale variations. The color index correlates with intensity during the rapid flares (the spectrum is flatter when the flux is higher), but it is rather insensitive to the long term trends. Over the 5 month observational period the light curve shows an overall brightening of about 1 mag followed by a fast decline. The EGRET pointing covers part of the very bright phase ($V \sim 13.2$) and the initial decline. An ultraviolet spectrum was also obtained with the International Ultraviolet Explorer (IUE) (1200-3000 Å) during the EGRET observations. The variability of the optical emission of S5 0716+714 by itself sets important constraints on the magnetic field strength and on the physical processes responsible for it. Interpreting the whole electromagnetic spectrum with synchrotron self Compton models leads to the prediction of a bright γ -ray state during the EGRET pointing. We discuss how the γ -ray data could be used as a diagnostic of the proposed models.¹

Key words: BL Lac objects: general – BL Lac objects: individual: S5 0716+714 – gamma-rays: theory – ultraviolet: galaxies

1. Introduction

One of the most important recent discoveries in the field of active galactic nuclei (AGN) is that flat spectrum radio sources (blazars) emit a substantial fraction, sometimes most of their power, at γ -ray energies. About 60 flat spectrum radio sources have been detected up to now above 100 MeV by the Energetic Gamma Ray Experiment Telescope (EGRET) onboard the Compton Gamma Ray Observatory (CGRO) (Fichtel 1994; von Montigny et al. 1995; Thompson et al. 1995). Strong and rapid variability seems to be a common property of these sources in the γ -ray band (e.g. von Montigny et al. 1995). The origin of this high energy emission is still unclear. A point of agreement is that the emitting plasma is in relativistic motion at small angles to the line of sight, with consequent beaming of the emitted radiation, which is required to avoid strong suppression of γ -rays due to photon-photon absorption (McBreen, 1979; for application to blazars see, e.g., Maraschi, Ghisellini & Celotti 1992; Dondi & Ghisellini 1995).

The correlated optical and γ -ray variability observed in some objects (Wagner et al. 1995, Hartman et al. 1996) indicates that the same population of electrons could be responsible for the flux in both bands, and suggests that the optical emission is due to the synchrotron process, while the γ -rays are produced by the inverse Compton process between highly relativistic electrons and soft seed photons. However, the origin of the seed photons, the location and size of the emitting region(s), and the degree of relativistic beaming of the high energy radiation are unknown.

Send offprint requests to: G. Ghisellini

^{*} Table 2 is only available in electronic form at CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsweb.u-strasbg.fr/Abstract.html>

¹ Table 2 is only available in electronic form at CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsweb.u-strasbg.fr>

According to one interpretation, the seed photons can be produced by synchrotron emission internally to the emitting blob or jet (Maraschi, Ghisellini & Celotti 1992). Another possibility is that the blob itself could illuminate a portion of the broad line clouds, whose reprocessed line radiation can dominate the radiation energy density in the blob (Ghisellini & Madau 1996). In another scenario the seed photons are produced externally, by the accretion disk (Dermer & Schlickeiser 1993), by the broad line region illuminated by the disk (Sikora, Begelman & Rees 1994), and/or by some scattering material surrounding the jet (Blandford 1993, Blandford & Levinson 1995). Finally, a dusty torus could provide IR radiation to be scattered in the γ -ray band, as suggested by Wagner et al. (1995).

For plausible values of the parameters, the relativistic electrons which emit γ -rays by the inverse Compton process emit synchrotron photons in the IR-optical-UV range. As a consequence, the two emissions are expected to be strongly correlated. Optical observations can therefore be crucial to understand the nature of γ -ray emission in blazars since they allow to derive information on both the seed photons and the radiating high energy electrons.

Renewed interest in optical monitoring is motivated also by the discovery that in the radio domain the flux of some blazars varies by 10-30% on a timescale less than a day (Quirrenbach et al. 1992). The inferred brightness temperature, of the order of 10^{18} K, poses severe difficulties to theoretical models, since it cannot be reconciled with the Compton limit of 10^{12} K for the commonly estimated values of the beaming factor (see, e.g. Blandford, 1990). Possible explanations are that these variations are due to regions with extremely high beaming or to coherent processes or that they are extrinsic, due to interstellar scintillation. Simultaneous optical and radio monitoring is thus critical to discriminate different models.

In this paper we present the results of recent optical-UV observations of the BL Lac object S5 0716+714. This source is an ideal target for both the research fields outlined above, since it has been detected several times in the γ -ray band and has shown intraday variability at optical and radio wavelengths (Wagner et al. 1996). The optical observations were carried out between Nov. 15, 1994 and Apr. 30, 1995, the International Ultraviolet Explorer (IUE) pointings were on Feb. 27 and March 1, 1995 and cover the EGRET-CGRO viewing period (1995 Feb. 14-28).

The outline of the paper is as follows: in Sect. 2 we review the available information on S5 0716+714 at all wavelengths; in Sects. 3 and 4 we describe and discuss the optical observations, taken from different sites and telescopes, and the UV observation obtained with IUE simultaneously with the phase 4 EGRET pointing; in Sect. 5 we model the overall emission spectrum with a homogeneous synchrotron self Compton (SSC) model and with a relativistic inhomogeneous jet model. In Sect. 6 we draw our conclusions.

2. The BL Lacertae object S5 0716+714

S5 0716+714 was discovered in the Bonn-NRAO radio survey (Kühr et al. 1981) of flat spectrum radio sources with a 5 GHz

flux greater than 1 Jy, and was identified as a BL Lacertae object, because of the featureless spectrum, by Biermann et al. (1981). From the stellar appearance on deep plates Stickel, Fried & Kühr (1993) derived a lower limit of $z > 0.2$ on the redshift, while Schalinski et al. (1992) give $z > 0.3$. Proper motions of VLBI components have been observed (Schalinski et al. 1992), and would result in an apparent superluminal motion with $\beta_{app} = 4.6$ if $z = 0.3$ and $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

A high dynamic range VLA 1.5 GHz radio map shows a double structure (Antonucci et al. 1986) resembling that of an FR II radiogalaxy seen end-on, but higher frequency (5 and 8.5 GHz) VLA maps do not show the double structure (Wagner et al. 1996). The FR II morphology conflicts with the suggestion that the parent population of BL Lac objects should be FR I radiogalaxies (see e.g. Urry & Padovani 1995 for a review) but it is consistent with a more general “unified” scheme whereby both FR I and FR II radiogalaxies have similar relativistic jets in their cores (Maraschi & Rovetti 1995).

A further “anomaly” is that the observed ratio of the core to the extended radio flux, which is $R = 4.7$ at 5 GHz and $R \sim 1$ at 1.5 GHz (Perley, Fomalont & Johnston 1982), is at the low extreme of the R -distribution of the radio-selected BL Lac objects (see e.g. Ghisellini et al. 1993 and references therein).

The radio core has an inverted spectrum [$\alpha = -0.42$ between 1.5 and 5 GHz; Perley, Fomalont & Johnston 1980; $F(\nu) \propto \nu^{-\alpha}$] The overall flat radio spectrum seems the result of the superposition of the core and more extended and steeper components (see the decomposition of the radio spectrum in Eckart et al. 1986), even if the intraday radio variability suggests that at least some of the radio emission may come from much smaller regions.

The source was monitored in the mm band by Steppe et al. (1992, 1993) and Reich et al. (1993). At these frequencies (90-230 GHz) the flux varied by a factor 2-3 on timescales of the order of a month. As shown by Reich et al. (1993), the radio flux at all the monitored frequencies increased after Jan. 1992, the time of the first detection by EGRET-CGRO.

In the far infrared it was detected by the Infrared Astronomical Satellite (IRAS) at 12, 60, and 100 μm (Impey & Neugebauer 1988).

Despite S5 0716+714 is a very interesting source it has been very poorly observed in the past. Biermann et al. (1981) report the two photographic magnitudes 13.2 and 15.5, the former derived from the POSS plate. The two previous visual magnitude estimates of 11 (Kühr et al. 1981) and 19 ± 1 (Perley, Fomalont & Johnston 1980) cannot be considered as fully reliable because no clear source for them is specified by these authors. The 8-month long relative R light curve, from September 1990 to March 1991 (Heidt & Wagner 1996, Wagner et al. 1996), shows a variation amplitude of about 2.5 mag. No historic curve, based on archive plates, to search for possible long term flux changes, is available. Takalo, Sillanpää & Nilsson (1994) found a high level of optical polarization, which varied from 20% to $\sim 16\%$ in two days.

In X-rays the source was observed by HEAO1, EINSTEIN, and ROSAT (Biermann et al. 1981, 1992; Wagner, 1992; Cappi

et al. 1994; Comastri, Molendi & Ghisellini 1995). During the ROSAT observations the flux varied by 70% in 1000 seconds and by a factor 7 in 2 days (Wagner 1992, Cappi et al. 1994). EGRET observed and detected S5 0716+714 several times from January 1992 (Lin et al. 1995; von Montigny et al. 1995). The integrated flux above 100 MeV varied between 1.3 ± 0.5 and $5.3 \pm 1.3 \times 10^{-7}$ photons $\text{cm}^{-2} \text{s}^{-1}$. Low statistics did not allow to derive a significant change of the energy spectral index during variations of the flux (Lin et al. 1995). Combining observations of four different viewing periods, Lin et al. (1995) found an average energy spectral index $\alpha_\gamma = 0.85 \pm 0.20$.

Following the discovery of intraday radio variability with a $\sim 10\%$ amplitude (Heeschen et al. 1987), the source was monitored several times at radio wavelengths, always showing intraday variability (Quirrenbach et al. 1989, 1992). At optical frequencies, variations on timescale of ~ 1 day with amplitude of several tenths of magnitude are reported by Wagner et al. (1996).

A simultaneous optical (R band) and radio (5 GHz) monitoring campaign, performed for a month (February 1990), suggested that the variations in the two bands were correlated, with amplitudes up to a factor 3 in the optical and 25% in the radio, and with typical timescales of 1 and 7 days in both bands (Wagner et al. 1990; Quirrenbach et al. 1991). Further support to the suggestion of a correlated optical and radio variability comes from the correlation between the radio spectral index in the range 5-8.4 GHz and the optical flux (Wagner et al. 1996 and references therein). If the redshift $z > 0.3$ the implied brightness temperature exceeds 10^{18} K (Quirrenbach et al. 1991).

3. Observations and data analysis

3.1. Optical

Photometry was carried out with several telescopes: the 1.05 m astrometric telescope of the Observatory of Torino (B and R bands mainly), the 0.4 m automatic telescope of the University of Perugia (Tosti, Pascolini & Fiorucci 1996), the 0.5 m telescope of the Astronomical Station of Vallinfreda, and the 0.4 m telescope of the IAS-CNR at the Observatory of Monte Porzio (the last two operated by the Rome group). All telescopes were equipped with CCD cameras and filters B , V (Johnson) and R , I (Cousins). Data reduction was performed with different software packages: the standard IRAF and MIDAS procedures and some others developed locally, as the ROBIN program developed at the Observatory of Torino by L. Lanteri (private communication). All images were bias and flat-field corrected. The comparison among data obtained in the same nights with different telescopes yielded the same results within the errors, which are typically 0.02 mag in R and 0.03-0.05 mag in B .

We did not find any published comparison sequence in the field of S5 0716+714; only one nearby star was calibrated by Takalo, Sillanpää & Nilsson (1994). We therefore calibrated some reference stars using several Landolt's (1992) fields. The stars are indicated in the map of Fig. 1 and the corresponding magnitudes are listed in Table 1; quoted errors include the uncer-

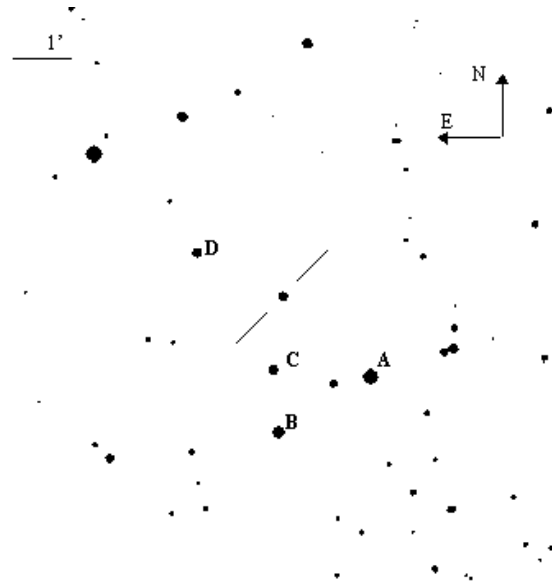


Fig. 1. Finding chart of S5 0716+714. See Table 1 for the magnitudes of the labeled reference stars.

Table 1. Reference stars in the field of S5 0716+714

	I	R	V	B
A	10.92 ± 0.04	11.21 ± 0.04	11.51 ± 0.04	12.07 ± 0.06
B	11.79 ± 0.05	12.12 ± 0.05	12.48 ± 0.05	13.06 ± 0.08
C	12.85 ± 0.05	13.18 ± 0.05	13.58 ± 0.06	14.17 ± 0.08
D	12.97 ± 0.04	13.27 ± 0.04	13.66 ± 0.04	14.25 ± 0.05

tainty on the absolute calibration, while magnitude differences between the reference stars in a given filter are good within 0.02 magnitudes. The star C of this sequence is the one already measured by Takalo, Sillanpää & Nilsson (1994): our magnitudes are consistent with theirs.

The reddening correction in the direction of S5 0716+714 was taken $A_V = 0.23$. It was derived from the N_H value of $3.9 \times 10^{20} \text{ cm}^{-2}$ (Ciliegi, Bassani & Caroli 1993), in agreement with the best fit values of both UV and X-ray observations (Cappi et al. 1994), using the conversion formula by Shull & van Steenberg (1985) and $R=3.09$ (Riecke & Lebofsky 1985). From the curve of Riecke & Lebofsky (1985) we derived also $A_B = 0.30$, $A_R = 0.19$, and $A_I = 0.14$. In the following we use these values when deriving fluxes from the apparent magnitudes.

The observation journal is reported in Table 2: here, for each date are given the measured magnitudes and their uncertainties in the four filters together with the used telescope. In some occasions several images were made in the course of the same night to search for intraday variability; these nights are indicated by an asterisk. The magnitude uncertainties have been evaluated by taking into account the mean dispersion of the internal error, but not the errors on the absolute calibration. The entire data set can be retrieved via anonymous ftp at the address [ftp.to.astro.it](ftp://ftp.to.astro.it/pub/blazars/) directory `/pub/blazars/`.

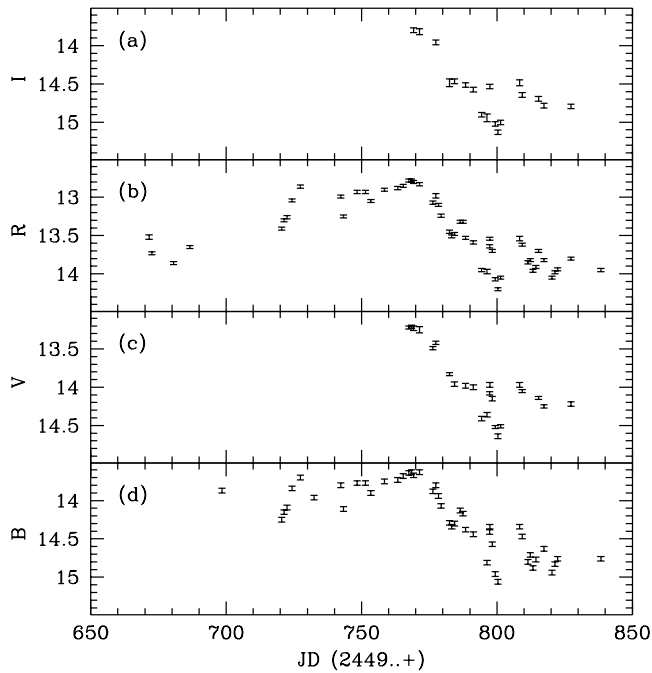


Fig. 2a–d. Light curve of S5 0716+714 in the I, R, V, B bands.

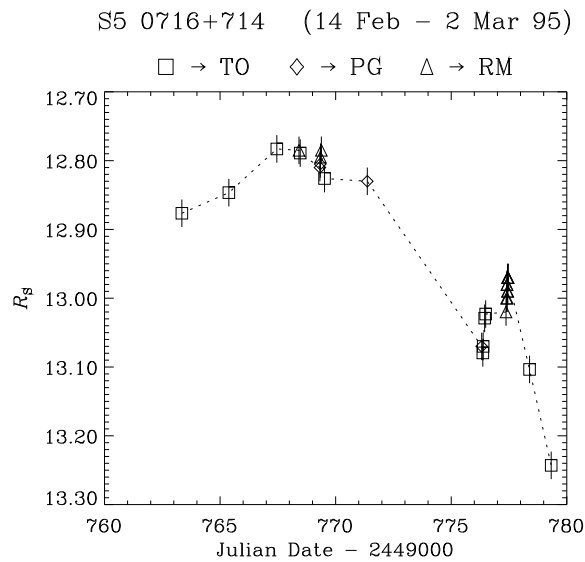


Fig. 3. Enlargement of the light curve in the R filter during the period of the EGRET pointing.

The light curves in the four bands are shown in Fig. 2a,d; in Fig. 3 the part of the R light curve (the most sampled in that period) corresponding to the EGRET pointing is shown in more detail.

3.2. Optical and ultraviolet spectra

Optical spectrum

S5 0716+714 was observed with the 1.5 m Cassini telescope of the Loiano Station, operated by the Observatory of Bologna.

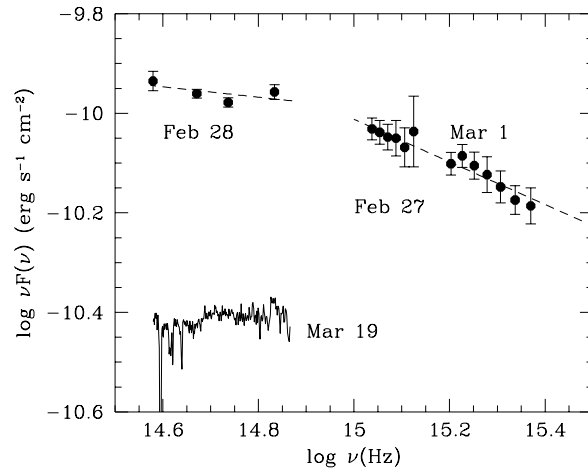


Fig. 4. The (binned) UV spectrum is plotted together with simultaneous optical data and with the optical spectrum made 19 days after, as labelled.

The 4000–8000 Å spectrum was obtained on March 19, 1995 at 00:58 UT, with the BFOSC all-transmitting spectrograph. A Thomson 1024 × 1024-pixel coated CCD detector was used with 300 l/mm grating and 2''.5 slit to give a spectral resolution of ~ 25 Å. The total integration time was 2400 s. The 2-dimensional frame was biased and flat-field corrected using standard techniques, and for the extraction of the one-dimensional spectrum the optimal extraction algorithm (Horne 1986) was adopted. The wavelength-calibration was performed with a fifth-order polynomial fitting to the dispersion curve of He–Ar comparison lamp spectrum. The spectrum has been flux-calibrated with the observation of the standard star BD +33 2642 (Stone 1977) and corrected with the local mean atmospheric extinction curve. The zero point of the spectral-flux calibration is not very accurate due to the poor seeing (~ 3 arcsec), and to the non-photometric conditions during the night.

The resulting spectrum is shown in Fig. 4, where no significant features are visible. The continuum spectrum was fitted with a power-law model assuming the reddening correction value $A_V = 0.23$ and the parametric formula for the mean extinction law computed by Cardelli et al. (1989), derived from the Rieke & Lebofsky (1985) extinction curve. The best fit spectral index and associated 1σ error is $\alpha = 0.95 \pm 0.05$.

Ultraviolet spectrum

S5 0716+714 was observed with the two cameras aboard IUE on February 27 (LWP) and on March 1, 1995 (SWP, see Table 3). The spectral signal was extracted from the line-by-line image files using the GEX routine (Urry & Reichert 1988) running within MIDAS, and flux calibrated using the curves provided by Bohlin et al. (1990, SWP) and Cassatella et al. (1992, LWP); the resulting spectral flux distribution, where no apparent emission or absorption features are visible, is also given in Fig. 4. The flux was integrated and averaged over bins of 100 Å, and the statistical uncertainties associated with the resulting flux points

were evaluated as in Falomo et al. (1993). To the flux errors a 1.25% photometric error was added, according to Edelson et al. (1992). The SWP and LWP flux distributions were separately fitted to a power-law model taking into account the reddening correction $A_V = 0.23$, and the extinction curve of Cardelli et al. (1989). The fitted spectral parameters at the 1- σ confidence level are given in Table 3.

The optical photometry indicates that the source was stationary during the 2 days which separate the IUE observations, thus, assuming absence of significant interday UV variability a power-law fit has been attempted to the joined SWP and LWP spectral flux distributions (1230-2800 Å). The resulting spectral index is intermediate between those obtained from the individual fits ($\alpha_\nu = 1.43 \pm 0.03$), with an acceptable value of χ_r^2 . This, together with the fact that the slope of the high energy part of the UV spectrum is not significantly larger than that of the lower energy region ($\Delta\alpha \sim 0.2$), excludes the measurable presence of a curvature or of a bump in the spectrum at ~ 2000 Å.

A power-law fit to the merged SWP and LWP spectra with the interstellar extinction taken as a free parameter yields a minimum χ^2 for A_V coinciding with the galactic value. Including the simultaneous optical data (see the B , V , R and I data points in Fig. 4) the combined optical+UV fit yields $\alpha = 1.27 \pm 0.02$, with $\chi_r^2 = 5.1$ (16 d.o.f.). The BVRI data are fitted by a power law of slope $\alpha = 1.1 \pm 0.1$, significantly flatter than the UV slope, therefore one can infer that there is a break in the spectrum which occurs around the U band.

We finally note that the flux at 1400 Å ($F_\nu = 3.21 \pm 0.07$ mJy), is more than 20 times larger than that measured in 1980 (IUE exposure SWP 9440; Pian & Treves 1993).

4. Results of the optical monitoring

4.1. The light curves

From the light curves of Fig. 2a,d, one can see that the source brightened after JD=2449720, remained bright for ~ 50 days, and then declined rapidly, to reach a minimum around JD=2449800. Note that in the ‘high’ state the source almost reached its recorded historical maximum.

Superposed to this relatively long term behaviour, there are fast fluctuations, whose amplitude appears wider in the ‘low state’, indicating that they could contain the same amount of energy as in the ‘high state’. To make this behaviour more apparent we modelled the ‘long’ trend by interpolating a smooth curve through the local minima of the R light curve. We used a cubic spline interpolation, which is plotted as dashed line in Fig. 5a. We stress that this spline interpolation crucially depends on the chosen minima of the light curve, and therefore it should be considered only as a possible representation of the source behaviour. The percent excesses between the observed fluxes and those derived from this interpolating curve are shown in Fig. 5b: labels from A to E correspond to five events of ‘short’ timescale. The event C is the best sampled, whereas the shape of the event B can depend on the adopted spline and could be considered as the superposition of two (or even more) events.

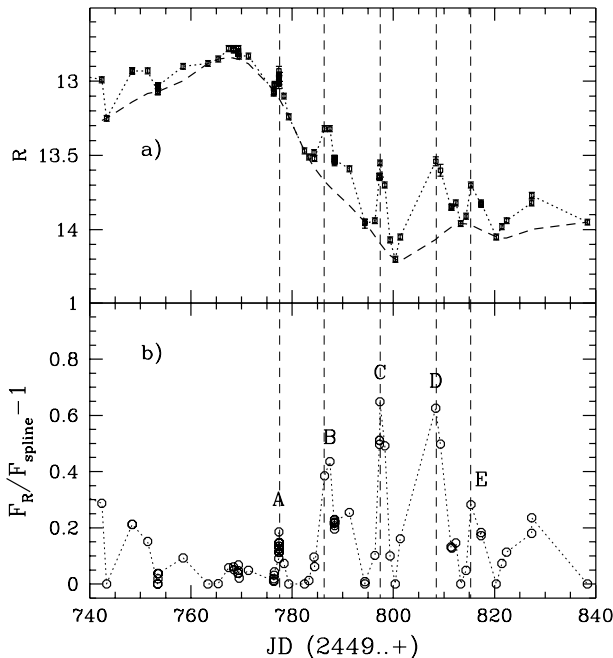


Fig. 5a and b. The light curve in the R band (panel a) and the adopted spline function passing through the local minima of the light curve (dashed line). In panel b) we plot the ratio (minus 1) between the observed flux and the spline value. The labels and the vertical dashed lines identify interesting flaring events.

Notice that the event A has the shortest duration and the smallest variability amplitude. For the best sampled peak in this curve, the rising and decaying times appear approximately equal, with no indications of a ‘plateau’ longer than one day.

4.2. Intranight variations

Rapid variations of the flux of S5 0716+714 have been observed in different occasions (e.g. Wagner & Witzel, 1995). In our campaign we also searched for intranight variations especially in the R band (see Table 2). Variations with an amplitude of about 5% in few hours in the course of the same night were observed in some occasions. Some of the resulting light curves in the B and R bands are shown in Figs. 6-7. The lower panels in these figures show the difference between the instrumental magnitudes of the two reference stars used in the night (normalized to the mean value).

Notice that in the two nights of 1995 Jan. 9 (Fig. 6) and Feb. 4 (Fig. 7) evident trends in the source brightness are apparent, like if our sampling was limited to a fraction of a longer time scale change. These cases, however, are important because the probability of their occurrence can be evaluated by means of the *run test* - see, for instance, the introductory textbook by Barlow (1989) - which is non-parametric and independent of the measured uncertainties. In the night of Jan 9, 1995, we have 9 data points in the B band and other 10 in the R band, both showing decreasing brightness trends and having each one two runs only. The expected run numbers for these two data samples

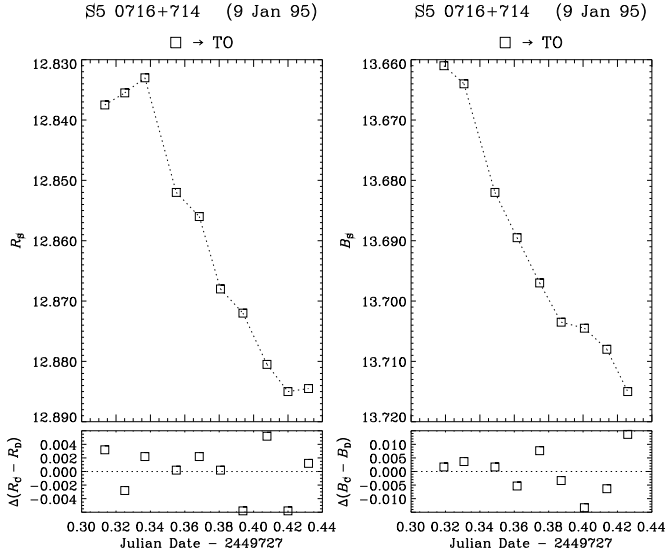
Table 2. IUE observations

	$\Delta\lambda$	λ (Å)	α_ν	$F(\nu)$ (mJy)	χ^2/n	d.o.f.
SWP 54002	1230-1930	1600	1.60 ± 0.09	3.98 ± 0.07	0.99	6
LWP 30122	2200-2800	2600	1.39 ± 0.23	7.87 ± 0.14	0.25	5
SWP+LWP	1230-2800	2000	1.43 ± 0.03	5.45 ± 0.08	1.01	12

$\Delta\lambda$ indicates the wavelength range used in the analysis.

1 σ errors are given both for the spectral index and for the flux.

The monochromatic flux $F(\nu)$ is calculated at the frequency corresponding to the wavelength λ .

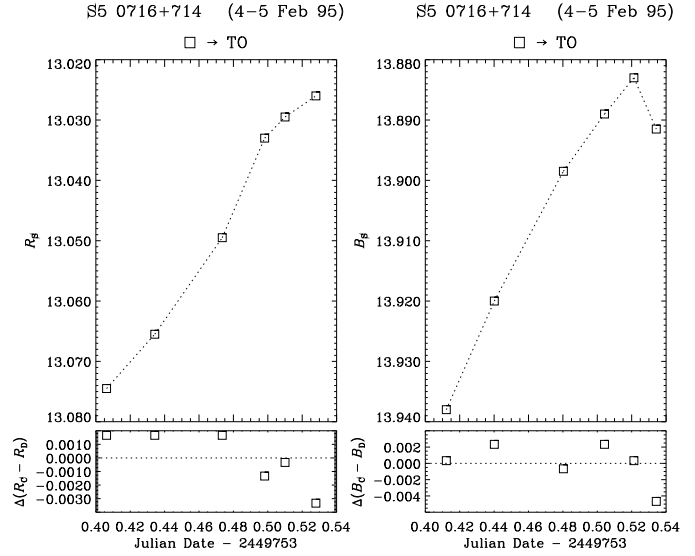
**Fig. 6.** Intranight variations during the Jan. 9 1995 night, in the R band (left panel) and in the B band (right panel)

are 5.44 (B band) and 6.0 (R band), corresponding to 2.5 and 2.7 standard deviations, respectively. Their probabilities, estimated from a one-tailed gaussian distribution are 6.2×10^{-3} and 3.7×10^{-3} ; the combined chance probability to observe these trends at the same time and in the same direction in the two bands can be estimated as 1/2 times their product, because the observations in the two bands can be considered independent, as indicated by the comparison star behaviour. We obtain then the value of 1.2×10^{-5} , which gives a high statistical significance to this effect.

Again on Feb 4-5 a well defined trend in the same direction is apparent in the two bands. The run test gives a total probability of 5.9×10^{-4} , greater than the previous value because of the smaller number of data points, but statistically significant.

4.3. Flux-spectral index correlation

We computed the spectral index α (by a fit procedure), after dereddening the observed fluxes, for the nights in which at least three filter data are available, including intranight data. These nights cover the period from February 18 to the beginning of April in which the source brightness declined by 1.4 magnitudes. The resulting values of α vary between 0.81 and 1.15, with a mean value of 0.94, and do not have any significant correlation with the flux: the linear correlation coefficient with the

**Fig. 7.** Intranight variations during the Feb 4 1995 night, in the R band (left panel) and in the B band (right panel)

V magnitude is only 0.15. In particular the spectral slopes measured when the source was at the maximum (February 19) and minimum (March 23) brightness were both equal to 0.99.

We also computed the 2-point spectral index α_{BR} , which can be compared with the spectral index calculated using more filters. The differences between these indices are usually less than the statistical uncertainties and this gives confidence that α_{BR} traces the ‘true’ optical spectral index. Again these values vary in the range 0.83-1.2.

We found that when the source reached the ‘low’ state (after JD=2449790) the spectral index responded to the fast flux changes in the sense that the spectrum was flatter when the source was brighter. This is evident from Fig. 8 where R and α_{BR} are plotted against time. The linear correlation coefficients of α_{BR} with B and R were 0.627 and 0.477, respectively, and the corresponding chance probabilities are 7.1×10^{-3} and 5.3×10^{-2} . As pointed out by Massaro & Trevese (1996) the latter estimate is biased towards a lower significance of the correlation and therefore this probability must be considered as a conservative indication of the correct significance. A less biased estimate of the true correlation coefficient can be obtained by taking the mean of the two previous values, and in this case the corresponding probability results equal to 2.1×10^{-2} . We can conclude that S5 0716+714 shows a significant ($\geq 95\%$) pos-

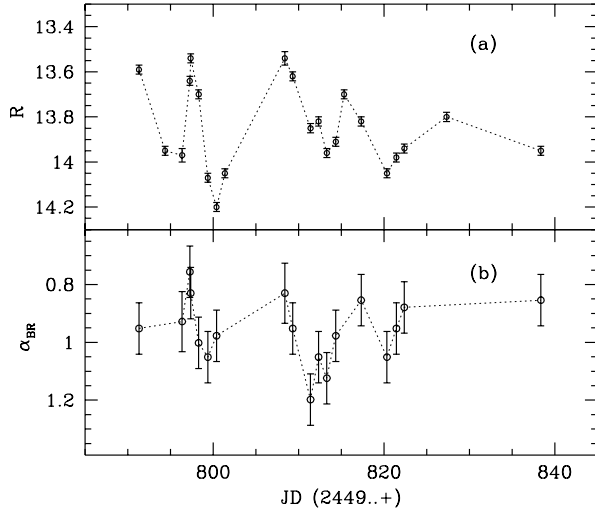


Fig. 8a and b. The light curve in the R band during the ‘low’ state of S5 0716+714 (panel a) compared with the corresponding spectral index α_{BR} (panel b).

itive correlation between the spectral index and the flux only during rapidly varying changes.

5. Discussion

5.1. Short time variations

The structure of the light curves of S5 0716+714 together with the color index behaviour suggest that *two* different processes are operating in the source: the first one is responsible for the long term “achromatic” rise and decay of the brightness, while the second one is responsible for the fast flares for which the color index (α_{BR}) is significantly bluer at higher intensity. It is possible to interpret the first behaviour in terms of emission produced in a large region, in which a stable energy injection mechanism operates, at least on few month timescale.

The fact that the rising and decaying timescales of the short term flares are comparable, with much shorter plateaux, bears important consequences. In fact, it is highly unlikely that these two timescales correspond to the particle acceleration and cooling timescales, respectively, since a priori there is no reason for them to be equal, and the observed equality would be coincidental. Instead, the equality is naturally explained in two frames:

1) As Wagner et al. (1993) suggested for the similar behaviour of S4 0954+658, the characteristic variability pattern of the flares could be caused by a curved trajectory of the relativistic emitting blob; in this case the viewing angle θ (i.e. the angle between the blob velocity vector and the line of sight) varies with time and consequently the beaming or Doppler factor $\delta = [\Gamma(1 - \beta \cos \theta)]^{-1}$ changes too. Since the observed monochromatic flux can be assumed to be proportional to $\delta^{3+\alpha}$, symmetric θ variations can account for the observed flare shape. In this case, for rather small amplitude changes, one has

$$\frac{\Delta F}{F} \simeq (3 + \alpha) \frac{\Delta \delta}{\delta} \simeq (3 + \alpha) \beta \Gamma \delta \Delta \theta \quad (1)$$

which shows that, for $\Gamma \sim \delta \sim 1/\sin \theta \sim 10$, a flux variation of 20% - see the results of Sect. 5.2 - implies a change of θ of less than 0.5 degrees only.

2) The rising and decay times are both associated with the light crossing time. This can happen if both the electron injection and cooling processes operate on time-scales shorter than R/c (where R is the source radius), since in this case the light crossing time is the only characteristic time that can be observed. On the other hand, the short duration of a plateau also requires that either the acceleration or the cooling times (or both) are not much less than R/c . Assume in fact that *both* these times are much shorter than R/c . Then at any given time the received flux originates in a different partial volume of the source. Different times correspond to different sub-volumes of the source and, if these regions are equal, the observed flare would have a plateau.

The situation is further complicated by the contribution of the steadier component, which dilutes the flux and spectral index variations.

We can envisage a test to discriminate between the two possibilities. In case 1) the variable component should have a constant spectral index: amplification due to different degree of beaming would in fact be ‘achromatic’ (if there is no spectral break within the considered bands). The sum of the variable (flatter) and the steady (steeper) components results in a continuous flattening of the spectrum as the source brightens. The flattest spectral index should correspond to the maximum flux. In case 2) the variable component has instead a variable spectral index: the maximum flux corresponds to integration over the source volume, whose sub-regions contain electrons of different ages. Therefore the maximum flux *does not* correspond to the flattest spectral index.

Our data are not sufficient to discriminate between the two possibilities, since a better sampling (in more bands, including IR and, possibly, UV) is needed. Nevertheless, if possibility 2) turns out to be preferred by future investigations, it will be possible to draw important consequences about the value of the magnetic field strength and the particle energy. We will derive these two parameters in the next section by independent arguments, and it is interesting to see if they agree.

Assume therefore that the rapid flares are due to quasi-instantaneous injection of electrons and rapid cooling. In particular we require that the synchrotron and self Compton cooling times are equal to or shorter than R/c :

$$t_{cool} = \frac{6\pi mc^2}{\sigma_T c \gamma_o B^2 (1 + U_S/U_B)} \leq \frac{R}{c}, \quad (2)$$

where σ_T is the Thomson scattering cross section, $\gamma_o mc^2$ is the energy of the particles emitting in the optical, $U_B = B^2/(8\pi)$ is the magnetic energy density, and U_S is the radiation energy density of the synchrotron emission (as measured in the comoving frame). The observed frequency is given by

$$\nu_o = \frac{4}{3} \gamma_o^2 \nu_B \frac{\delta}{1+z}, \quad (3)$$

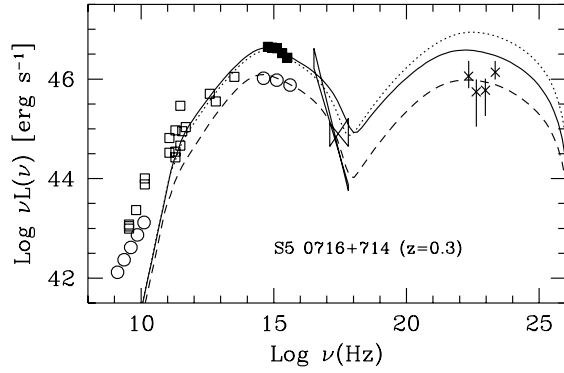


Fig. 9. Overall spectrum of S5 0716+714. Empty circles are from Wagner et al. (1996), ROSAT X-ray data are from Cappi et al. (1994) and γ -ray data are from Lin et al. (1995). Filled squares are optical and UV data from our observations, simultaneous (Feb 27 - Mar 1, 1995) with EGRET observations (not public yet). Empty squares are not simultaneous data from the literature. Lines refer to homogeneous SSC models, as discussed in the text.

where $\nu_B = eB/(2\pi mc)$ is the cyclotron frequency. By relating the radius R of the source with the observed variability timescale through $R \sim ct_{var}\delta/(1+z)$, we obtain the two limits:

$$B \geq 0.67 \nu_{15}^{-1/3} t_d^{-2/3} (1 + U_S/U_B)^{-2/3} \left(\frac{1+z}{\delta} \right)^{1/3}, \quad (4)$$

$$\gamma_o \leq 2 \times 10^4 \nu_{15}^{2/3} \left[\frac{t_d(1 + U_S/U_B)(1+z)}{\delta} \right]^{1/3}, \quad (5)$$

where B is in Gauss, t_d is the variability timescale measured in days, and $\nu_o = 10^{15} \nu_{15}$ Hz. Note the relatively weak dependence on the beaming factor δ . Assuming $\delta/(1+z) = 10$, $t_d = 1$ day and $U_S/U_B = 1$ (as indicated by the ratio between the self-Compton and synchrotron luminosities), we find $B \geq 0.2$ G and $\gamma_o \leq 10^4$.

5.2. Homogeneous synchrotron self Compton models

In Fig. 9 we show the overall spectrum of S5 0716+714 from radio to γ -rays, calculated assuming a redshift $z = 0.3$. Solid symbols in the optical (B , V , and R) and UV bands correspond to observations simultaneous (Feb. 27 and Mar. 1, 1995) with those of the CGRO pointing, whose data are not yet public. Empty squares are not simultaneous data taken from the literature. Empty circles refer to simultaneous radio to UV data published by Wagner et al. (1996), the X-ray ROSAT spectrum is from Cappi et al. (1994) and the γ -ray data are from Lin et al. (1995).

The spectrum of S5 0716+714, like that of other γ -ray loud blazars, is characterized by two broad peaks in ν - $\nu F(\nu)$ plots, which can be interpreted as due to the maximum output of the synchrotron and Compton components. Note that the knowledge of both the optical and UV spectra is crucial to define the frequency of the synchrotron peak.

Among the various ‘‘Compton’’ models (see Introduction) we focus on synchrotron self Compton (SSC) models, since no emission line has been observed in this source. This limits the importance of photons produced externally to the jet or emitting blob, which could serve as seed photons to be Comptonized in the γ -ray band.

Consider a spherical blob, having the variability timescale radius $R = ct_{var}\delta/(1+z)$ (in the comoving frame), with a tangled magnetic field B , emitting synchrotron and self Compton radiation. For appropriate particle distribution (to be better specified later), the synchrotron and self Compton spectra peak at the observed frequencies ν_S ($\sim \nu_o$), given by Eq. (3), and $\nu_C = (4/3)\gamma_o^2\nu_S$. From the ratio ν_C/ν_S we derive

$$B\delta \simeq \frac{\nu_S^2}{2.8 \times 10^6 \nu_C} (1+z) \quad (6)$$

The ratio of the luminosities of the two peaks measures the relative importance of the synchrotron and self-Compton powers. Considering the Compton scattering process occurring entirely in the Thomson regime, this ratio is directly related to the ratio between the magnetic and radiation energy densities of the synchrotron photons:

$$\frac{L_S}{L_C} \simeq \frac{U_B}{U_S} = \frac{4\pi R^2 c U_B \delta^4}{L_S} \quad (7)$$

Then

$$B\delta^3 \simeq (1+z) \left(\frac{2L_S^2}{L_C c^3 t_{var}^2} \right)^{1/2} \quad (8)$$

Equations (6) and (8) can be solved, yielding

$$\delta \simeq 1.67 \times 10^4 \left(\frac{\nu_C}{t_d \nu_S^2} \right)^{1/2} \left(\frac{L_{S,45}^2}{L_{C,45}} \right)^{1/4} \quad (9)$$

$$B \simeq 2.14 \times 10^{-11} (1+z) \frac{\nu_S^3 t_d^{1/2}}{\nu_C^{3/2}} \left(\frac{L_{C,45}}{L_{S,45}^2} \right)^{1/4} \quad (10)$$

where $L_{S,45}$ and $L_{C,45}$ are measured in units of 10^{45} erg s $^{-1}$ and t_d is t_{var} in days. Despite their simplicity, Eqs. (9) and (10) are a very powerful tool for deriving B and δ , but we must stress that for a reliable estimate we need to determine with great accuracy the simultaneous values of ν_C and ν_S , together with an estimate of the variability timescales.

For S5 0716+714, in the high state we have $L_S = 2.2 \times 10^{47}$ erg s $^{-1}$, $L_C = 4.4 \times 10^{47}$ erg s $^{-1}$, $\nu_S = 10^{15}$ Hz, $\nu_C = 8 \times 10^{22}$ Hz, and $t_{var} = 2$ days. With these values we derive $B \simeq 0.5$ Gauss, $\delta \simeq 11$, and $\gamma_o \simeq 7.7 \times 10^3$. Notice that these B and γ_o values are in agreement with the limits obtained in Sect. 5.1.

To calculate the emission spectrum we need to know the energy distribution of the emitting electrons. In a simple model with continuous injection of $Q(\gamma) \propto \gamma^{-s}$ [cm $^{-3}$ s $^{-1}$] relativistic electrons throughout the source, the relativistic electron density $N(\gamma)$ can be obtained by solving the well known steady continuity equation, leaving as free parameters the injected power L_{inj} , the spectral index s of the injected distribution and the extremes of the energy range γ_{min} and γ_{max} of the injected electrons

(see Ghisellini 1989 for details). For $s = 2.6$, the resulting $N(\gamma)$ can be approximated with a double power law, with indices 2 and 3.6 below and above γ_{min} , respectively. Therefore we can identify γ_{min} as γ_0 . The emission spectra, computed for low (dashed line) and high (solid and dotted lines) luminosity states are shown in Fig. 9. The two states correspond to an increase of the injected power L_{inj} from 10^{43} erg s $^{-1}$ to 3.3×10^{43} erg s $^{-1}$. The values of the other model parameters were $R = 3 \times 10^{16}$ cm, $\delta = 12$, $\gamma_{min} = 3 \times 10^3$ (or 4×10^3 , dotted line), $\gamma_{max} = 10^5$, and $B = 0.48$ (or 0.84, solid line) G. The two high state spectra differ for the values of B and γ_{min} : the one represented by the solid line has an increased magnetic field to maintain the ratio U_S/U_B constant and close to unity, while in the other the magnetic field is assumed to remain constant, and consequently the ratio U_S/U_B increases.

These two choices, both reasonable from a physical point of view, yield γ -ray luminosities which are a factor of 3 to 10 greater than the EGRET luminosities of 1992-1993. Furthermore, the emission is expected to vary simultaneously at all frequencies from the mm band up to the GeV band, although the amplitude of the synchrotron variations can be smaller than the Compton one.

5.3. Inhomogeneous jet model

We adopt the model of Ghisellini & Maraschi (1989) (see also Ghisellini, Maraschi & Treves 1985; Maraschi, Ghisellini & Celotti 1992) in order to check if relaxing the hypothesis of homogeneity changes substantially the inferred physical parameters. In this jet model the density of the emitting particles and the magnetic field strength decrease along the jet as power law functions of the distance x from the jet apex. Relativistic particles are assumed to have the same power law distribution, all along the jet, but the high energy cut-off (and the maximum emitted frequencies) are function of x . Spectral break and different spectral indices in the radiation spectrum are due to the different superpositions of the local spectra. The inner part of the jet has a parabolic shape, with cross sectional radius $r \propto x^{1/2}$, while the outer part is conical ($r \propto x$). In addition, the emitting plasma flowing in the jet is assumed to accelerate outwards in the parabolic region, while it has a constant bulk velocity in the conical part. We also fixed the dependence of the magnetic field with the cross sectional radius r of the jet, setting it to be $B \propto r^{-1}$, and the shape of the particle distribution ($N(\gamma) \propto \gamma^{-2}$, resulting in a fixed local spectral index, $\alpha_0 = 0.5$). We furthermore required that the photon-photon opacity must not be important all along the jet and that the region producing most of the optical can vary in a ~ 1 day timescale. With all these constraints, the set of parameters which better interpolate the data is rather well defined. These are listed in Table 4, and correspond to the spectra shown in Fig. 10. We also plot, for the low state, the emission coming from different regions of the jet.

The bulk of the γ -ray emission (up to a few GeV) is produced where the paraboloidal and the conical part of the jet connect, at distance $x \sim 5 \times 10^{17}$ cm from the jet apex, where the cross sectional radius of the jet is $r \sim 2 \times 10^{16}$ cm. As explained

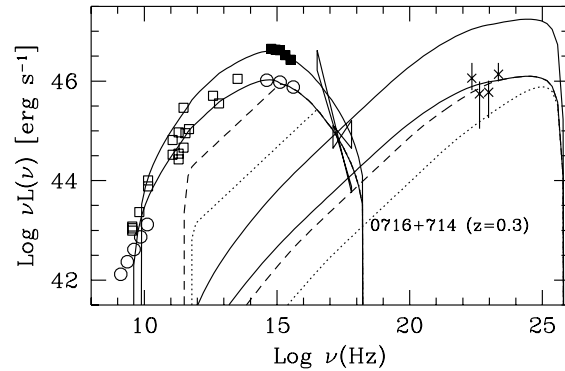


Fig. 10. The overall spectrum of S5 0716+714 interpreted as emission from an inhomogeneous relativistic jet. The solid lines refer to the emission from the entire jet, in the ‘high’ and ‘low’ state, according to the parameters listed in Table 4. The synchrotron and the self Compton components are shown separately. Dashed and dotted lines refer to the emission from regions within $x = 500x_0$ and within $x = 50x_0$ from the jet apex, respectively. See Table 4 for the definition of x_0 .

in Ghisellini, Maraschi & Treves (1985) and in Ghisellini & Maraschi (1989), this is expected in this class of models, due to the dependence of the Compton flux as a function of distance x , which is different in the parabolic and the cone zones. However, as can be seen also in Fig. 10, higher frequencies (both in the synchrotron and the self Compton branches) are produced closer in. This is not the consequence of the particular set of parameters, but is again a general characteristic of the model: requiring to fit the spectral change in the IR-UV-soft X-ray spectrum as a superposition of local spectra (each with a different high energy cut-off), inevitably implies higher frequencies be produced in the innermost regions of the jet. This characteristic can well be tested with future γ -ray satellite, such as GLAST, able to observe up to the 100 GeV band, where the emission is predicted to vary more dramatically and with the shortest timescales. Less dramatic differential variability is however expected also in the different EGRET energy bins. In the intermediate region between the paraboloid and the cone the magnetic field strength is $B \sim 0.5$ -1 G, the bulk Lorentz factor is $\Gamma \simeq 12$, the corresponding beaming factor is $\delta \simeq 14$, and the magnetic and radiation energy densities are close to equipartition.

In conclusion, we derive a set of parameters which agree well with the findings of the homogeneous model. With two important differences:

1) In the inhomogeneous jet model the synchrotron X-ray emission comes from the inner part of the paraboloid, at distances $x \sim 10^{15} - 10^{16}$ cm from the jet apex (and cross sectional radii $r \sim 10^{15} - 3.5 \times 10^{15}$ cm. Correspondingly, the X-rays should vary with the minimum timescale (the observed $t_{var} \sim 1$ hour), possibly explaining the very fast variability reported by Cappi et al. (1994).

2) In the inhomogeneous jet model variability is simultaneous only in selected bands. This can be seen in Fig. 10 where the dashed and dotted lines correspond to the synchrotron and the self-Compton emission arising in same regions of the jet:

Table 3. Input parameters for the inhomogeneous jet model. Note: $z = 0.3$ is assumed. For the parabolic part, the cross sectional radius r scales with distance x as $r \propto x^{1/2}$.

	Low state Paraboloid	Low state Cone	High state Paraboloid	High state Cone	Note
x_0 (cm)	10^{15}	5×10^{17}	10^{15}	5×10^{17}	initial distance
x_{max} (cm)	5×10^{17}	10^{19}	5×10^{17}	10^{19}	final distance
B_0 (Gauss)	12	0.54	12	0.54	initial magnetic field
N_0	3×10^5	2.8×10^3	6.7×10^5	6.3×10^3	initial electron density
α_0	0.5	0.5	0.5	0.5	local spectral index
m	1	1	1	1	$B \propto r^{-m}$
n	1.5	1.5	1.5	1.5	$N \propto r^{-n}$
a	0.4	0	0.4	0	$\Gamma \propto x^a$
γ_{max}	4.2×10^5	1.4×10^4	4.2×10^5	1.4×10^4	
Γ_0	1	12	1	12	initial bulk Lorentz factor
θ (degree)	4	4	4	4	viewing angle

~ 1 GeV emission should vary fast and simultaneously with 1-10 keV X-rays, while ~ 100 MeV γ -rays should vary simultaneously with the optical-UV emission and be characterized by a somewhat longer timescale. Spectral index variations in the IR-UV are possible if a perturbations travelling along the jet activates different regions, starting from the innermost one, responsible for the X-rays (for a detailed discussion, see Celotti, Maraschi & Treves, 1991). In this case the higher frequencies vary first, followed by the lower ones. Note that the predicted correlation between the flux and the spectral index is in the same sense as observed: flatter slopes when the flux is brighter.

Finally, note that in the inhomogeneous jet model, as well as in the homogeneous one, one cannot explain simultaneous intraday variability: even in the favorable case of a perturbation travelling down the jet, close to the line of sight, one should observe a time delay between the optical and the radio emission (even if shortened by relativistic effects).

6. Conclusions

Dense optical monitoring is a very powerful tool to better understand the radiation processes operating not only in the optical band, but also in the high energy part of the spectrum, and to put models for the γ -ray emission on test.

By themselves, the optical data we have obtained for S5 0716+714 are indicative of a complex behaviour possibly characterized by more than one component. The dependence of the spectral index on the fast variations of the flux and its insensitivity to longer-term trends, suggest localized injections of energy in the form of fresh and rapidly cooling relativistic particles, superimposed to a steadier emission, possibly coming from larger regions. Other constraints can be obtained by the shape of the fast flares, which seem symmetric and without plateau. In particular, this allows to derive a lower limit on the magnetic field strength. It is very interesting to verify if other sources behave like S5 0716+714 in this respect, to assess if this is a fundamental feature to be explained by emission models for blazars in general. Indeed, some of the best monitored sources (e.g. 3C 66A, Takalo et al. 1996; OJ 287, Sillampää et

al. 1996; 0954+658, Wagner et al. 1993; PKS 2155-304, Pian et al. 1997; PKS 0422+004, Massaro et al. 1996), seem to confirm this behaviour.

S5 0716+714 was very bright in the optical during the period of our observations, and in particular during the period of the EGRET pointing. Since electrons radiating optical synchrotron photons would emit γ -rays by the inverse Compton process, one inevitably predicts that the source had to be bright also in the EGRET band. The exact amount of the expected γ -ray emission is of course model-dependent (this is the reason why simultaneous observations are so important), but we can restrict the existing possibilities by noting that the absence of emission lines in the S5 0716+714 spectrum makes inverse Compton off photons produced *externally* to the jet less likely. For S5 0716+714, a simple homogeneous and one-zone model is appropriate to describe the optical/ γ -ray behaviour, because the regions emitting both these frequency bands are cospatial. This remains true in more realistic inhomogeneous models, which can better explain the entire electromagnetic output of the source, including the radio, far infrared, and X-ray bands. These models can be put on test by simultaneous monitoring in the above mentioned bands. In fact, a key ingredient of the inhomogeneous jet model we used is that smaller jet regions, located at short distances from the jet apex, are responsible for the high energy steep tail of the X-ray synchrotron spectrum, while larger and outer regions mainly contribute to the radio-to-IR band by synchrotron, and to hard X-rays by self-Compton emission. Perturbances travelling along the jet could be mapped in the frequency space, with shorter variability timescales corresponding to higher synchrotron and inverse Compton frequencies. In addition, there should be time delays in the light curves at different frequencies.

If the seed photons to be upscattered in energy are local synchrotron photons, then we predict that S5 0716+714 should have reached, during the February 1995 EGRET observation, its brightest ever detected γ -ray state, with a flux a factor 3-10 higher than in the 1992-1993 EGRET observing periods. Furthermore, the simple SSC models presented here assume that the bulk of the optical and γ -ray luminosities are produced

in the same region, and therefore the flux in both bands should vary together, without delay. In particular, we observed a decline of the optical flux on a time scale of a few days, and we expect that a similar trend should also be found in the γ -ray band if the statistical quality of the data is good enough. If the γ -ray emission is not linked with the decline of the bulk of the optical emission, one is still left with the interesting alternative that the γ -rays are connected with the optical fast flares. On the other hand, the absence of any optical- γ -ray correlation would be a severe problem for all simple SSC models.

References

- Antonucci, R.R.J., Hickson, P., Olszewski, E.W. & Miller, J.S., 1986, *AJ* 92, 1.
- Barlow, 1989, *Statistics*, J. Wiley & Sons, Chapter 8
- Biermann, P.L., et al. 1981, *ApJ* 247, L53
- Biermann, P.L., Schaaf, R., Pietsch, W., Schmutzler, T., Witzel, A. & Kühr, H., 1992, *A&AS* 96, 339.
- Blandford, R.D., 1990, in *Active Galactic Nuclei*, eds. T.J.L. Courvoisier, & M. Mayor (Berlin)
- Blandford, R.D. & Levinson, A., 1995, *ApJ* 441, 79
- Blandford, R.D., 1993, in *Compton Gamma Ray observatory*, ed. N. Gehrels & M. Friedlander (New York: American Institute of Physics), 533
- Bohlin, R., Harris, A.W., Holm, A.V., & Gry, C., 1990, *ApJS* 73, 413
- Cappi, M., Comastri, A., Molendi, S., Palumbo, G.G.C., Della Ceca, R. & Maccacaro, T., 1994, *MNRAS* 271, 438
- Cardelli, J.A., Clayton, G.C., Mathis, J.S. 1989, *ApJ* 345, 245.
- Cassatella, A., Gonzalez-Riestra, R., Imhoff, C., Oliverson, N., & Lloyd, C., 1992, *A&A* 256, 309
- Celotti, A., Maraschi, L. & Treves, A., 1991, *ApJ*, 377, 403
- Ciliegi, P., Bassani, L. & Caroli, E., 1993, *ApJSS* 85, 111
- Comastri, A., Molendi, S. & Ghisellini, G., 1995, *MNRAS* 277, 297
- Dermer, C. & Schlickeiser, R., 1993, *ApJ* 416, 458
- Dondi, L. & Ghisellini, G., 1995, *MNRAS* 273, 583
- Eckart, A., Witzel, A., Biermann, P., Johnston, K.J., Simon, R., Schalinski, C. & Kühr, H., 1986, *A&A* 168, 17
- Edelson, R., Pike, G. F., Saken, J. M., Kinney, A., & Shull, J. M., 1992 *ApJS*, 83, 1
- Fichtel, C.E. 1994, *ApJS* 90, 917
- Falomo, R., Treves, A., Chiappetti, L., Maraschi, L., Pian, E., & Tanzi, E. G. 1993, *ApJ* 402, 532
- Ghisellini, G., 1989, *MNRAS* 238, 449
- Ghisellini, G., Maraschi, L. & Treves, A., 1985, *A&A* 146, 204
- Ghisellini, G. & Maraschi, L., 1989, *ApJ* 340, 181
- Ghisellini, G., Padovani, P., Celotti, A. & Maraschi, L., 1993, *ApJ* 407, 65
- Ghisellini, G. & Madau, P., 1996, *MNRAS* 280, 67
- Hartman, R.C., Webb, J.R., Marscher, A.P., et al., 1996, *ApJ* 461, 698
- Heeschen, D.S., Krichbaum, T.P., Schalinski, C.J. & Witzel, A., 1987, *AJ* 94, 1493
- Heidt, J. & Wagner, S.J., 1996, *A&A* 305, 42
- Horne, K., 1986, *PASP* 98, 609
- Impey, C.D. & Neugebauer, G., 1988, *AJ* 95, 307
- Kühr, H., Pauliny-Toth, I.I.K., Witzel, A. & Schmidt, J., 1981, *AJ* 86, 854
- Kühr, H., Witzel, A., Pauliny-Toth, I.I.K. & Nauber, U., 1981, *A&AS* 45, 367
- Landolt, A.U., 1992, *AJ* 104, 340
- Lin, Y.C., Bertsch, D.L., Dingus, B.L., et al., 1995, *ApJ* 442, 96
- Maraschi, L., Ghisellini, G. & Celotti, A., 1992, *ApJ* 397, L5
- Maraschi, L. & Rovetti, F., 1995, *ApJ* 436, 79
- Massaro, E., Nesci, R., Maesano, M., et al., 1996, *A&A* 314, 87
- Massaro, E. & Treves, D., 1996, *A&A* 312, 810
- McBreen, B., 1979, *A&A* 71, L19
- Perley, R.A., Fomalont, E.B. & Johnston, K.J., 1980, *AJ* 85, 649
- Perley, R.A., Fomalont, E.B. & Johnston, K.J., 1982, *ApJ* 255, L93
- Pian, E. & Treves, A., 1993, *ApJ* 416, 130
- Pian, E., Urry, M.C., Treves, A., et al., 1997, *ApJ*, in press.
- Quirrenbach, A., Witzel, A., Krichbaum, T., Hummel, C.A., Alberdi, A. & Shalinski, C., 1989, *Nat* 337, 442
- Quirrenbach, A., Witzel, A., Wagner, S., et al., 1991, *ApJ* 372, L71
- Quirrenbach, A., Witzel, A., Krichbaum, T.P., et al., 1992, *A&A* 258, 279
- Reich, W., Steppe, H., Schlickeiser, R., Reich, P., Pohl, M., Reuter, H.P., Kanbach, G & Schönfelder, V., 1993, *A&A* 273, 65
- Riecke, G.H. & Lebofski, M.J., 1985, *ApJ* 288, 618
- Schalinski, C.J., Witzel, A., Krichbaum, T.P., Hummel, C.A., Quirrenbach, A. & Johnstone, K.J., 1992, in *Variability in Blazars*, ed. E. Valtaoja & M. Valtonen, (Cambridge Univ. Press), p.221
- Shull, J.M. & Van Steenberg, M.E., 1985, *ApJ* 294, 599
- Sikora, M., Begelman, & Rees, M.J., 1994, *ApJ* 421, 153
- Sillanpää, A., Takalo, L.O., Lehto, H.J., et al., 1996, *A&A*, 305, L17
- Steppe, H., Liechti, S., Mauersberg, R., Kömpe, C., Brunswig, W. & Ruiz-Moreno, M., 1992, *A&AS* 96, 441
- Steppe, H., Paubert, G., Sievers, A., et al., 1993, *A&AS* 102, 611
- Stickel, M., Fried, J.W. & Kühr, H., 1993, *A&AS* 98, 393
- Stickel, M., Padovani, P., Urry, C.M., Fried, J.W. & Kühr, H., 1991, *ApJ* 374, 431
- Stone, R.P.S., 1977, *ApJ* 218, 767
- Takalo, L.O., Sillanpää & Nilsson, K., 1994, *A&AS* 107, 497
- Takalo, L.O., Sillanpää, A. Letho, H.J., et al., 1996, *A&AS*, 120, 313
- Thompson, D.J., Bertsch, D.L., Esposito, J.A., et al., 1995, *ApJS* 101, 259
- Tosti, G., Pascolini, & Fiorucci, M., 1996, *PASP* 108, 706
- Urry, C.M. & Padovani, P., 1995, *PASP* 107, 803
- Urry, C.M., & Reichert, G. 1988, *IUE NASA Newsletter*, 34, 96
- von Montigny, C., Bertsch, D.L., Chiang, J., et al., 1995, *ApJ* 440, 525
- Wagner, S.J., Sanchez-Pons, F., Quirrenbach, A. & Witzel, A., 1990, *A&A* 235, L1
- Wagner, S.J., 1992, in "X-ray emission from AGN and CRB", MPE proceedings, p. 97
- Wagner, S.J., Witzel, A., Krichbaum, T.P. et al., 1993, *A&A* 271, 344
- Wagner, S.J. & Witzel, A., 1995, *ARA&A* 33, 163
- Wagner, S.J., Camenzind, M., Dreissigacker, O., et al., 1995, *A&A* 298, 688
- Wagner, S.J., Witzel, A., Heidt, J., et al., 1996, *AJ* 111, 2187